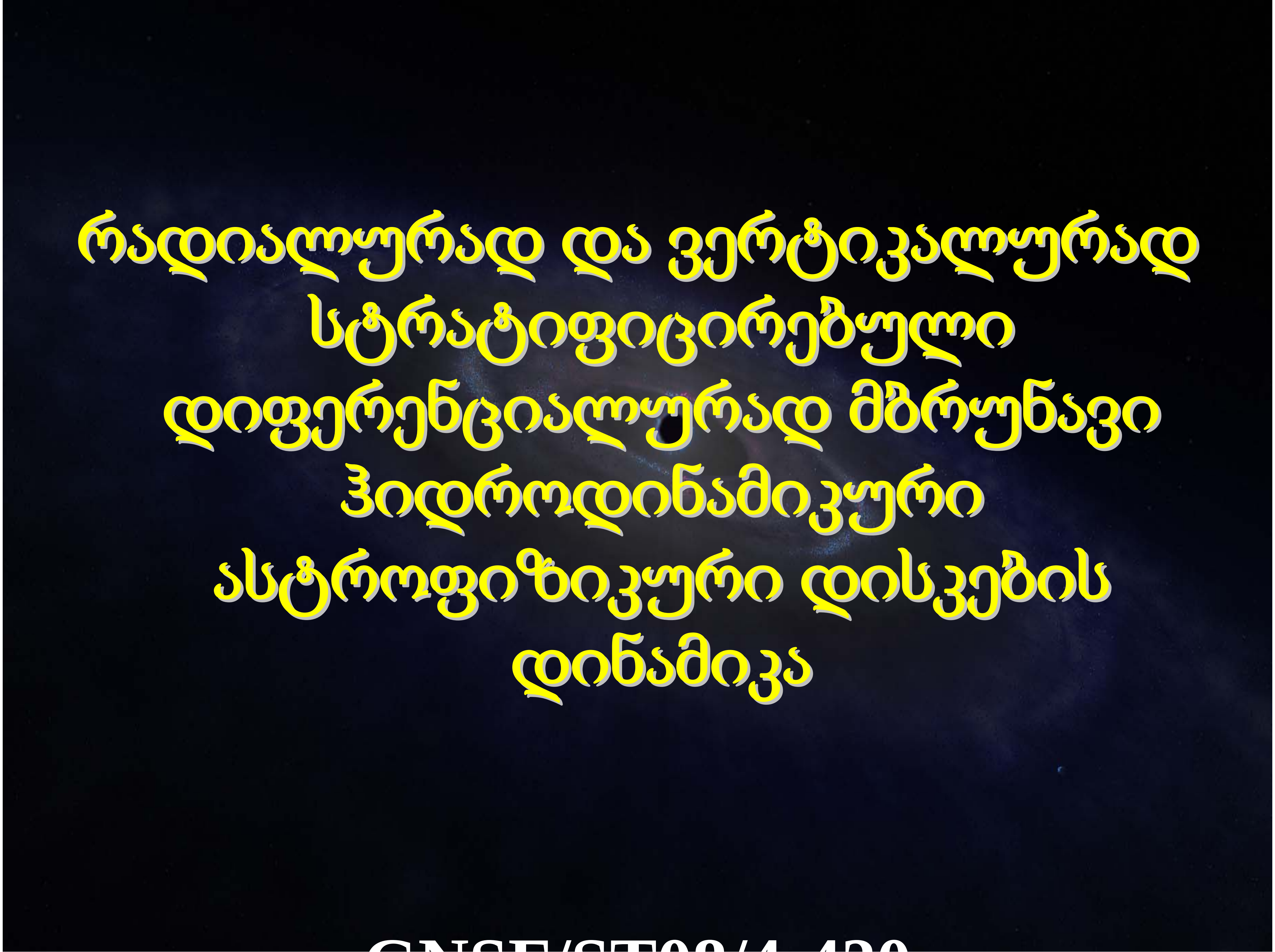


რადიალურად და ვერტიკალურად სტრატიფიცირებული დიფერენციალურად მბრუნავი ჰიდროდინამიკური ასტროფიზიკური დისკების დინამიკა

GNSF/ST08/4-420

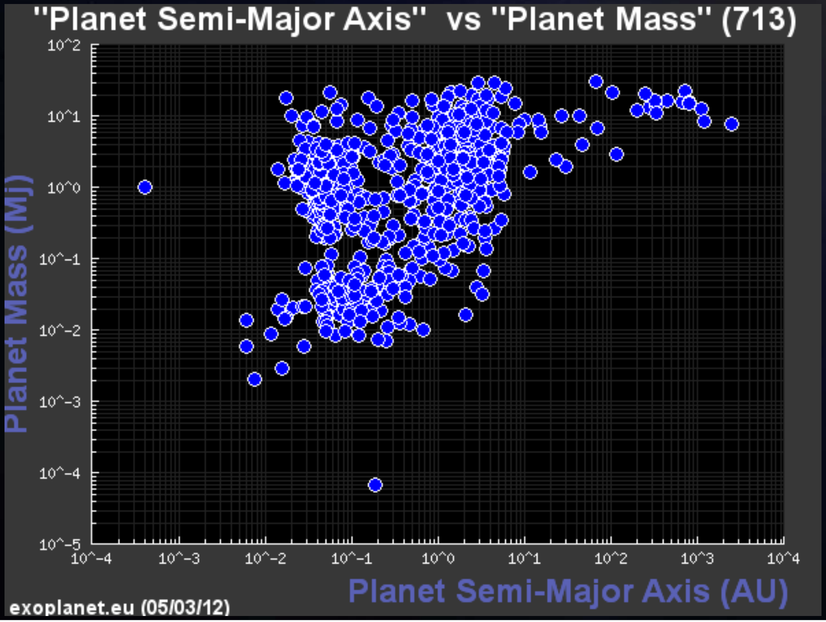
Exoplanets

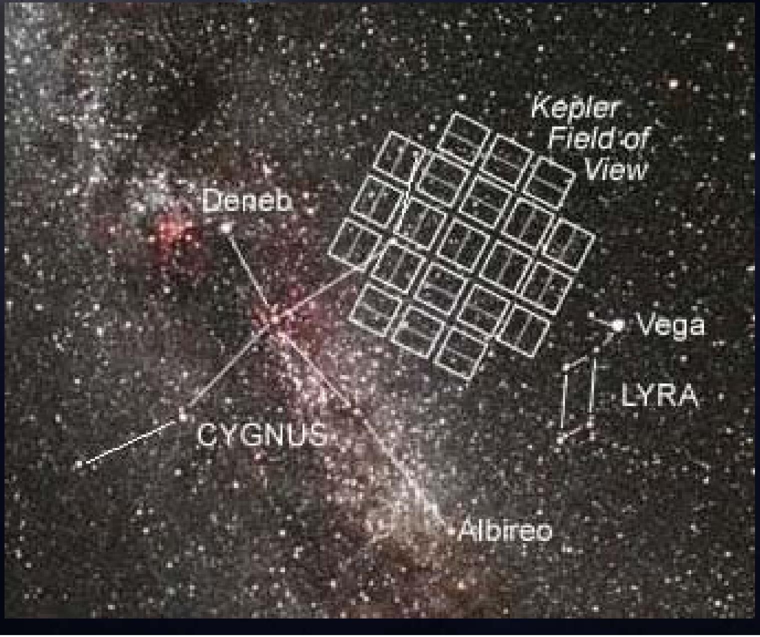
New era for the planet formation theories

Exoplanets found: 760

Exoplanets in multiple systems: 129

Kepler planet candidates: 2326

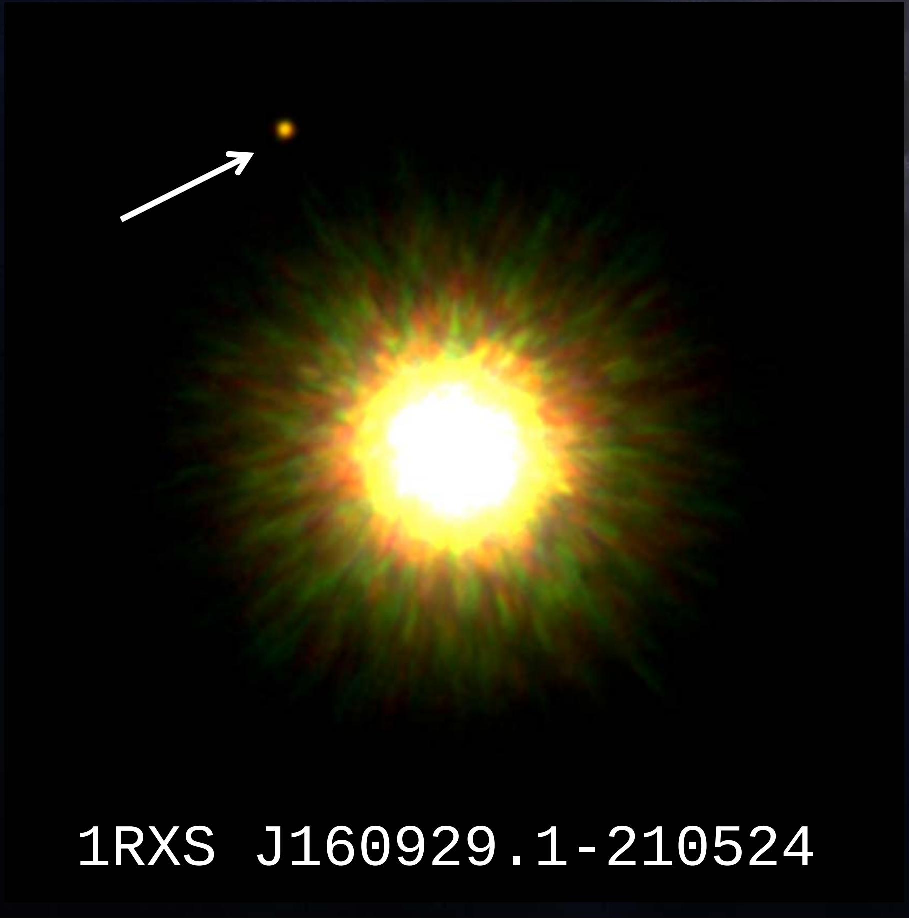




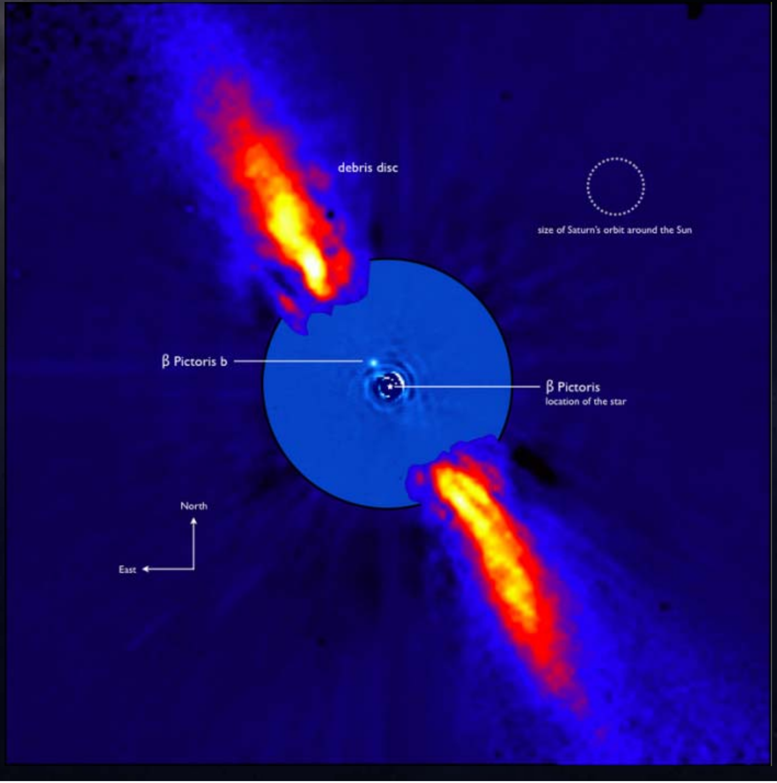
Observations of Planet Formation

Different stages of planet formation

Edge-On Protoplanetary Disk
Orion Nebula
HST • WFPC2
PRC95-45c • ST ScI OPO • November 20, 1995
M. J. McCaughrean (MPIA), C. R. O'Dell (Rice University), NASA

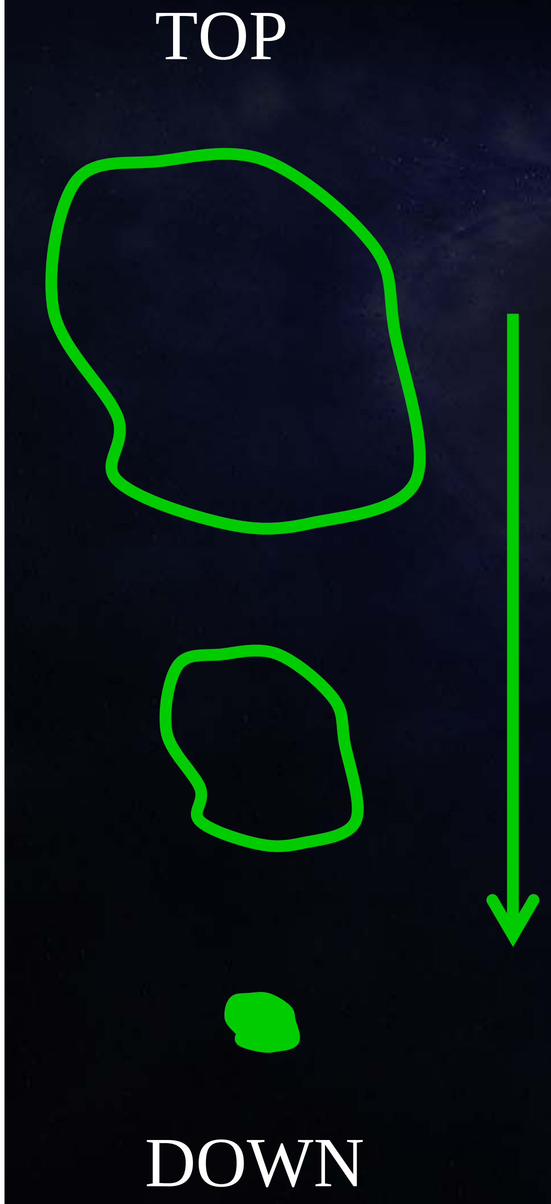


1RXS J160929.1-210524



Protoplanetary Disks

Planet Formation Theories



1. Top Down (*Laplace*)

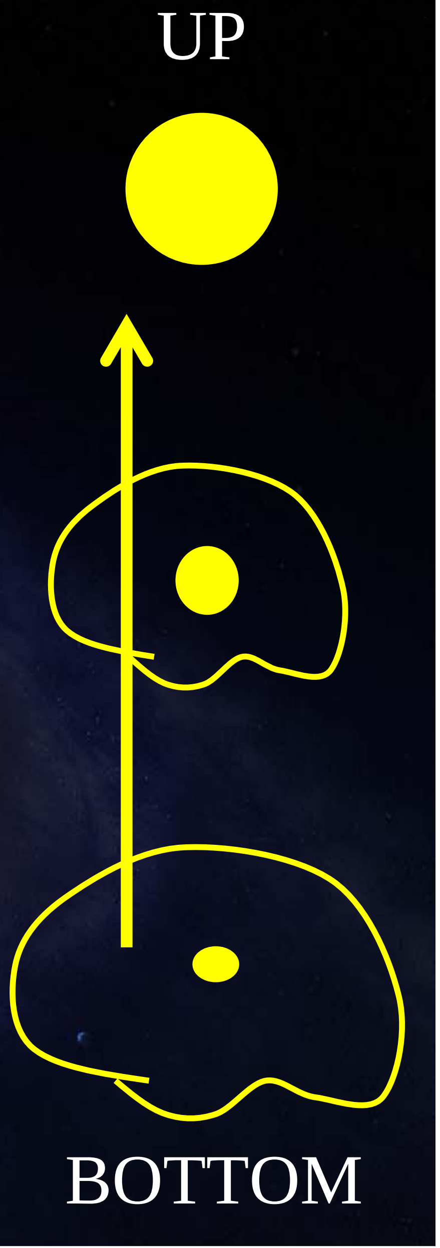
Gravitational

Fragmentation

2. Bottom Up (*Safronov*)

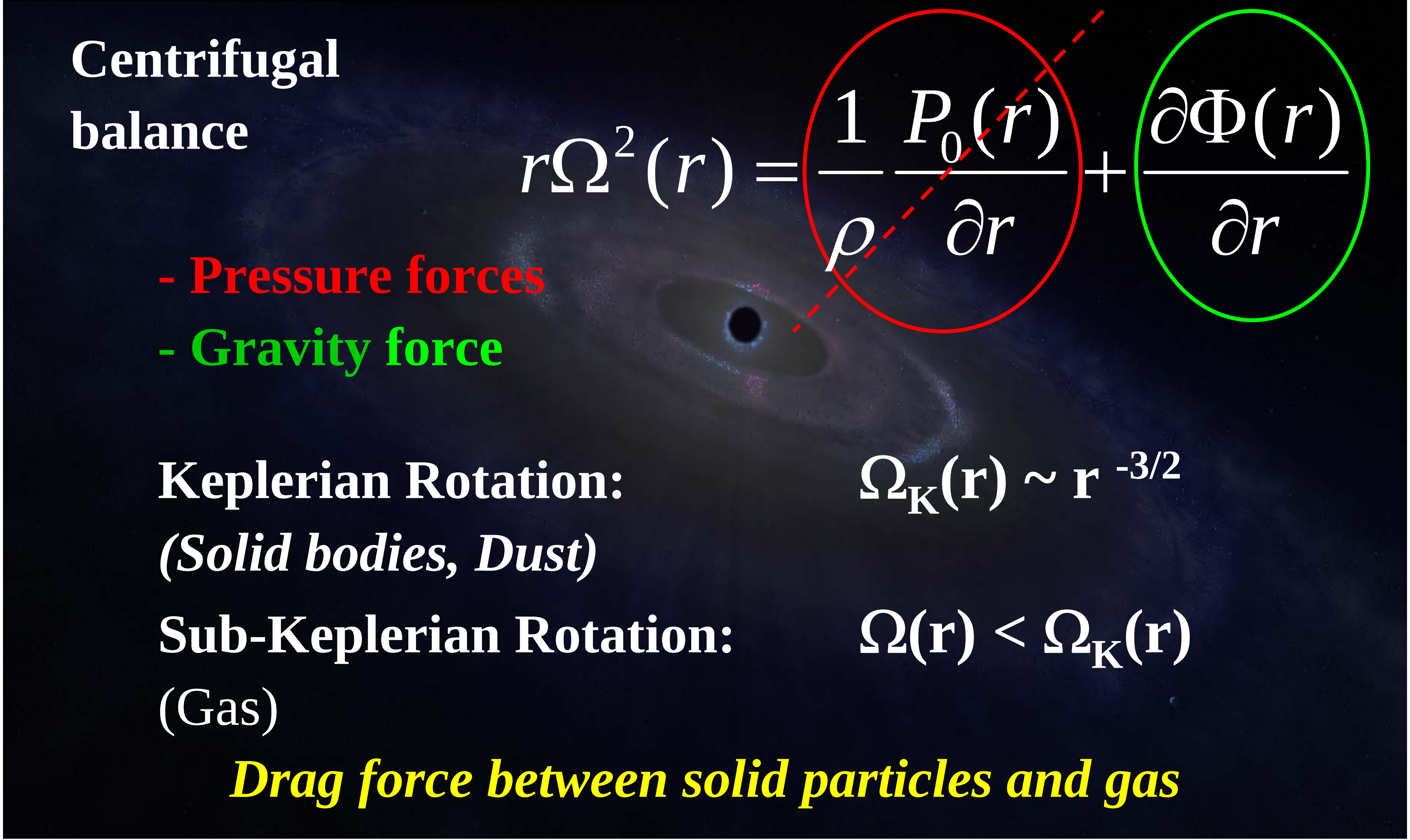
Core

Accretion



`dust2planet`

Keplerian Flows



Centrifugal balance

$$r\Omega^2(r) = \cancel{\frac{1}{\rho}\frac{P_0(r)}{\partial r}} + \frac{\partial\Phi(r)}{\partial r}$$

- **Pressure forces**
- **Gravity force**

Keplerian Rotation: $\Omega_K(r) \sim r^{-3/2}$

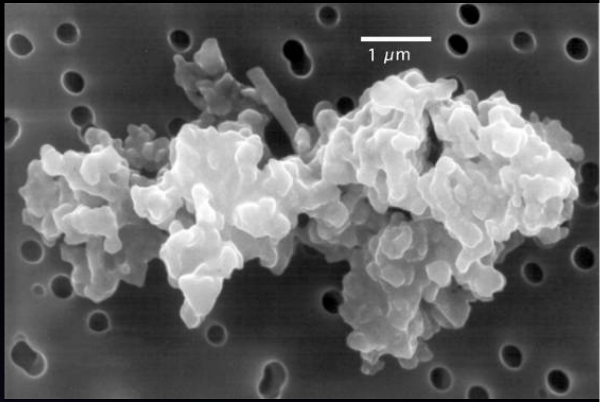
(Solid bodies, Dust)

Sub-Keplerian Rotation: $\Omega(r) < \Omega_K(r)$

(Gas)

Drag force between solid particles and gas

Dust Dynamics



Sub micro meter particles

growth by coagulation

Collisions: Sticking / Fragmentation

theory, simulations, experiment

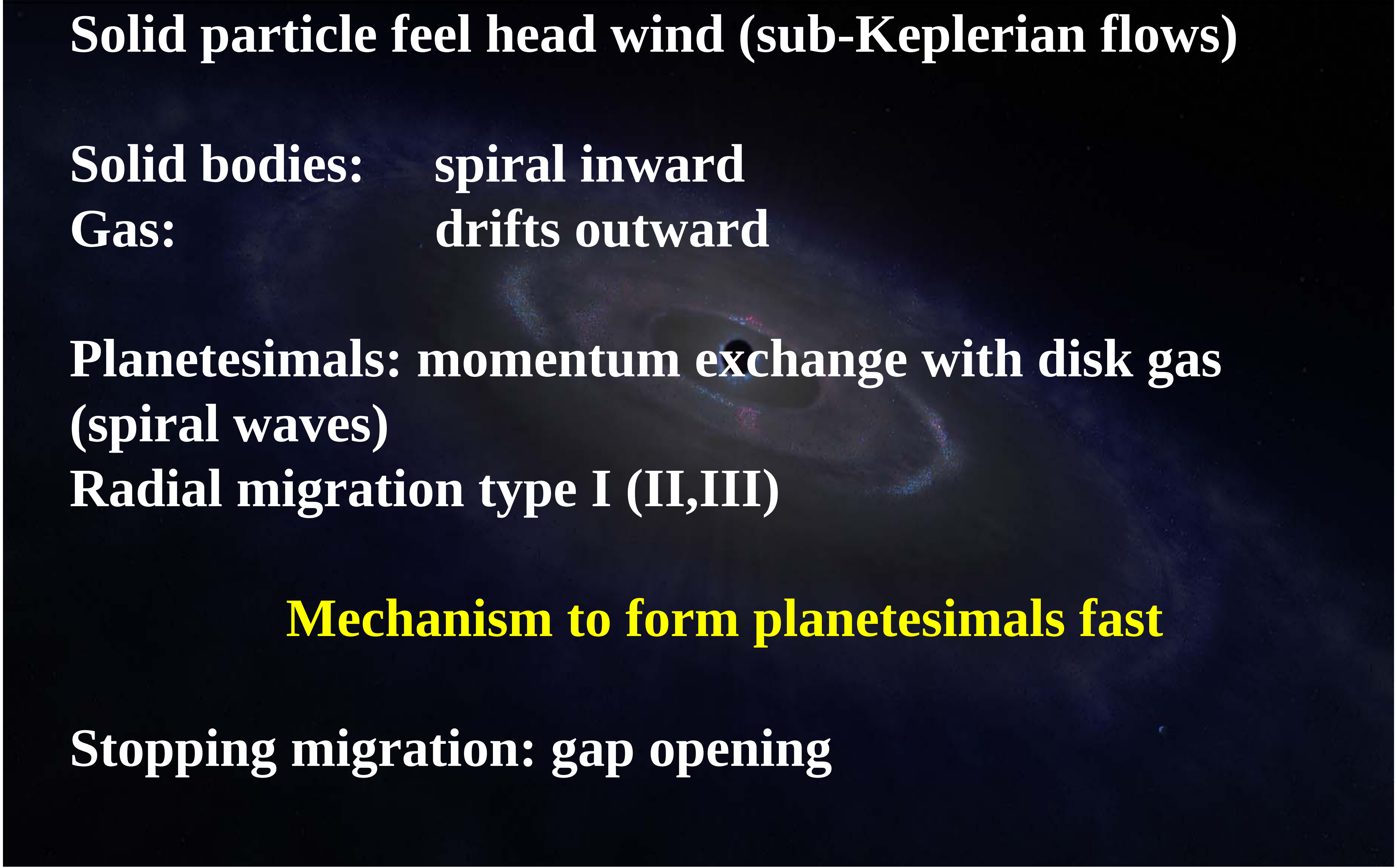
Fast coagulation: micro meter – cm

Meter to kilometers? SLOW

1 METER SIZE BARRIER

Planetesimals: ~km size bodies

Radial Migration



Solid particle feel head wind (sub-Keplerian flows)

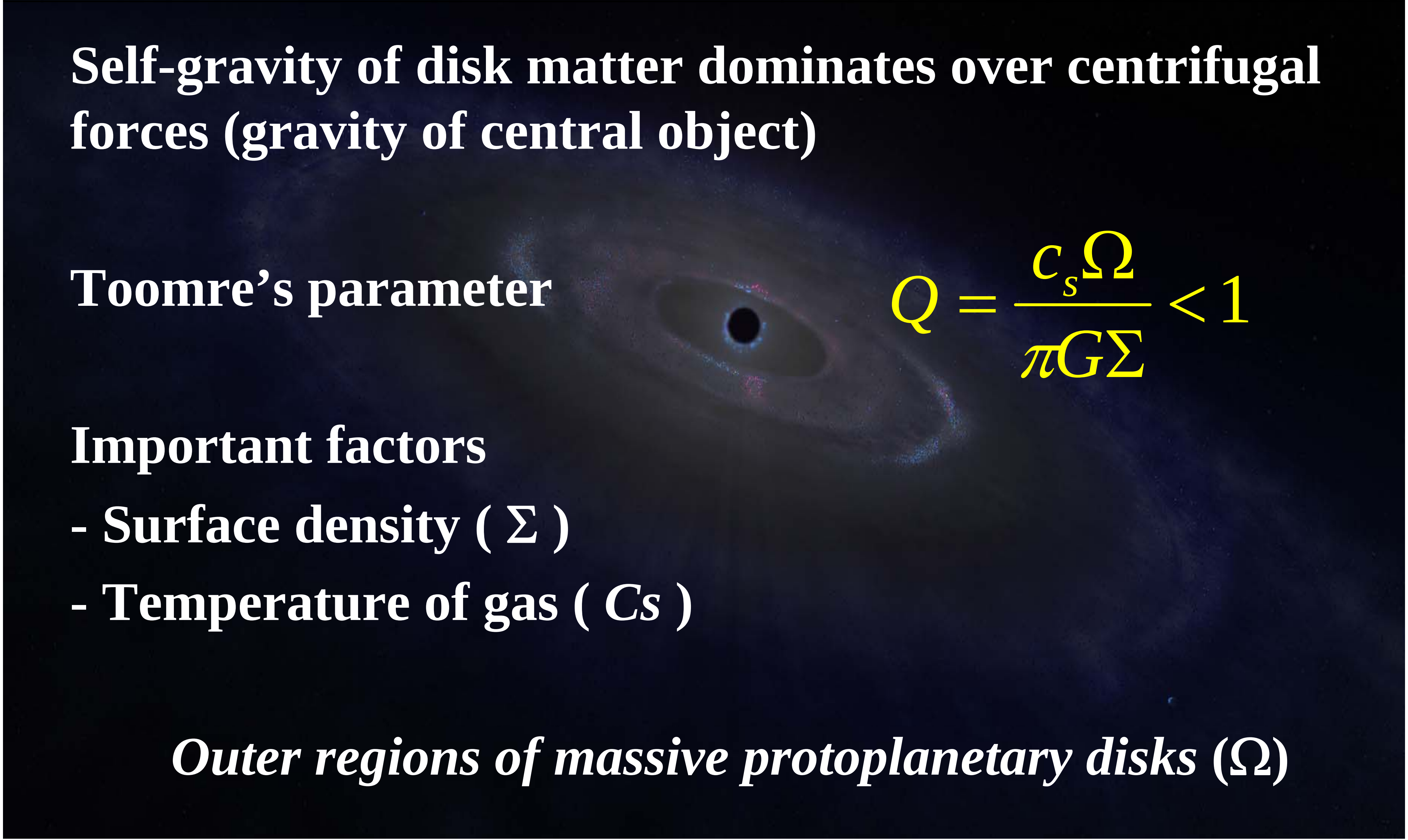
Solid bodies: spiral inward
Gas: drifts outward

Planetesimals: momentum exchange with disk gas (spiral waves)
Radial migration type I (II,III)

Mechanism to form planetesimals fast

Stopping migration: gap opening

Gravitational Instability



Self-gravity of disk matter dominates over centrifugal forces (gravity of central object)

Toomre's parameter

$$Q = \frac{c_s \Omega}{\pi G \Sigma} < 1$$

Important factors

- **Surface density (Σ)**
- **Temperature of gas (C_s)**

Outer regions of massive protoplanetary disks (Ω)

Gravitational Instability

Goldreich & Ward (1973)

Gas Disk + Dust sub-disk

Dust sedimentetion to central plane: gravitational instability

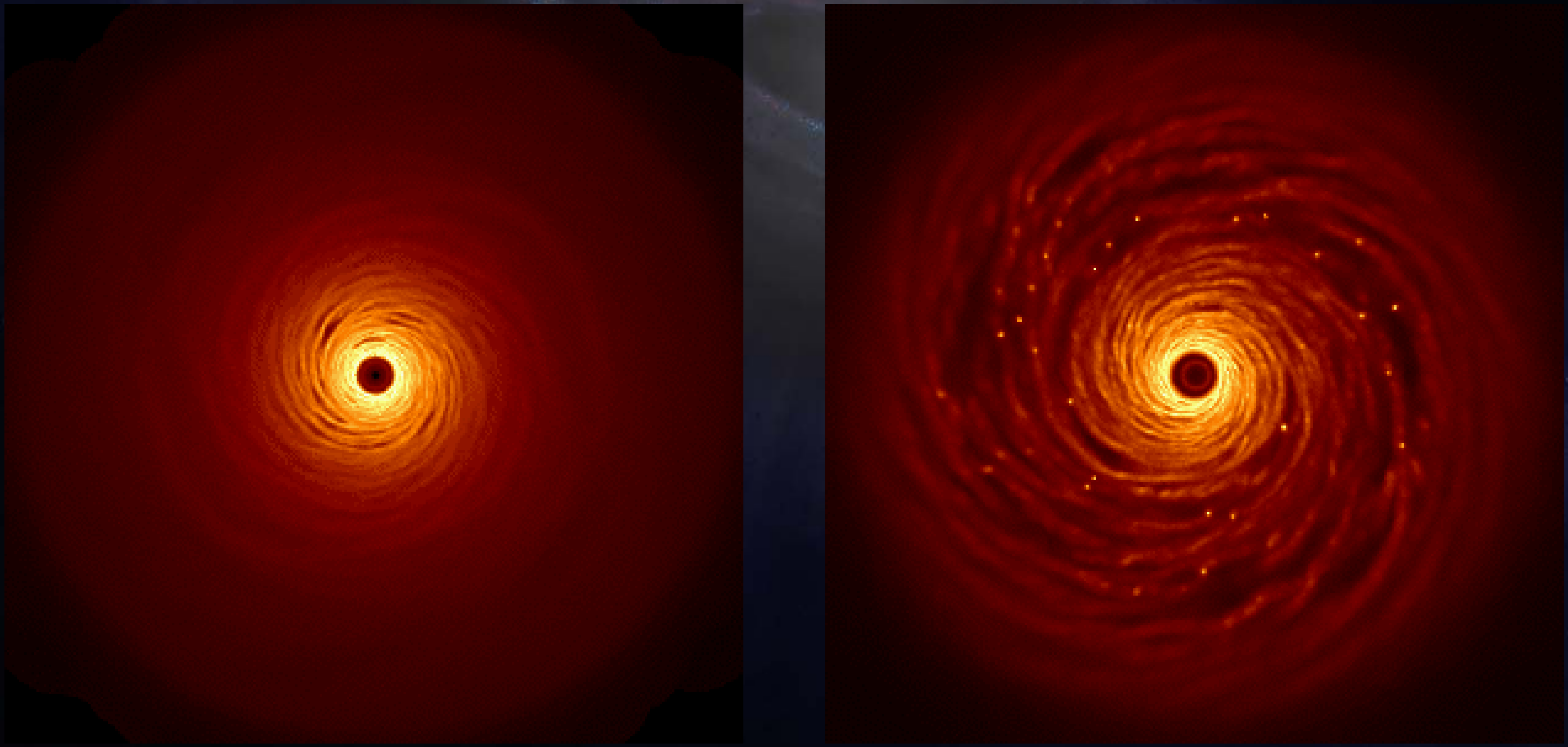
Direct formation mechanism

Triggering mechanism: Density-Spiral waves

Gravitational Instability

Fragmentation and formation of planetesimals

SPH simulations ***Rice et al. (2003)***



Gravitational Instability

Fragmentation: gas compression – heating

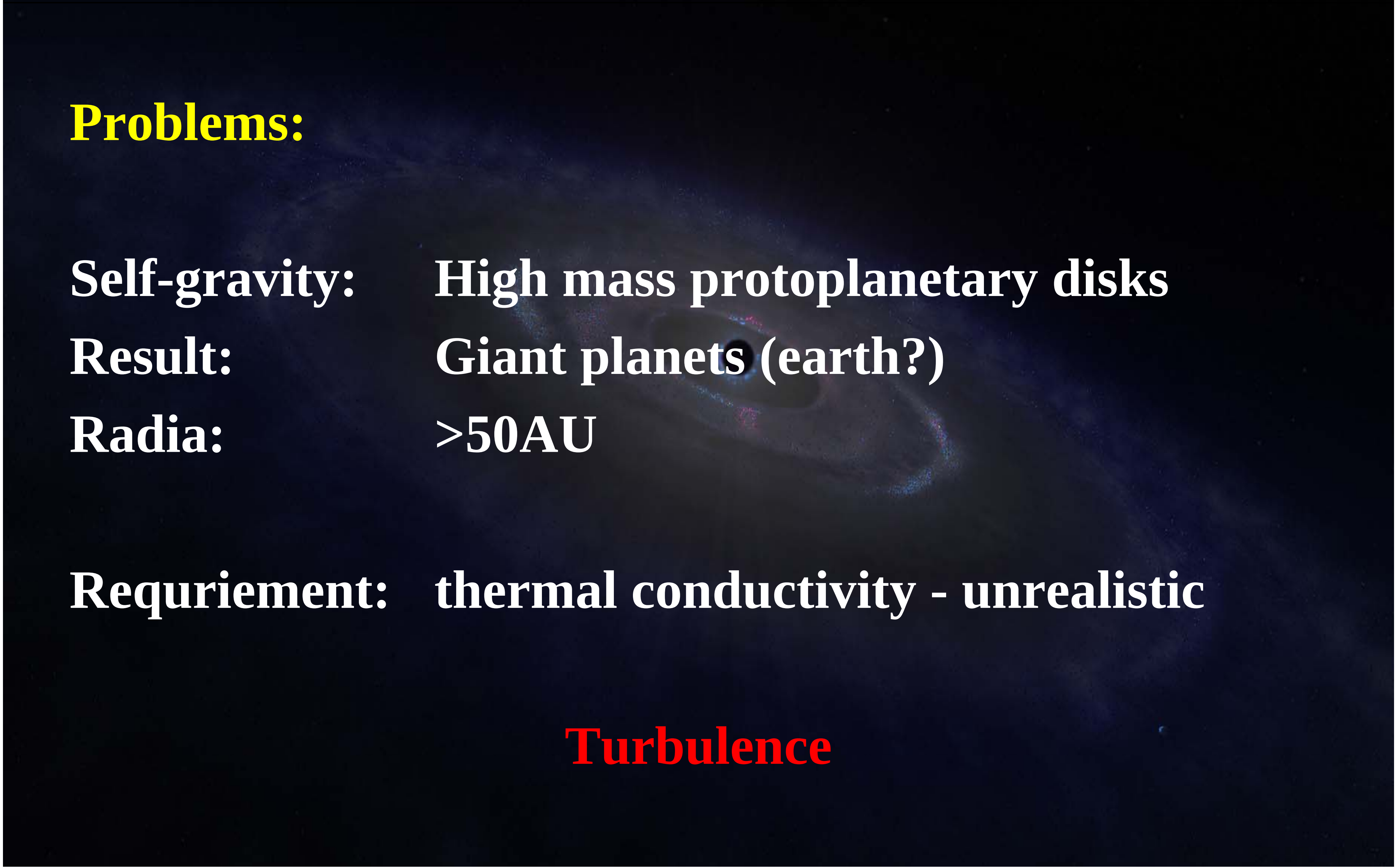
Instability: accelerated contraction due to self gravity;

Increasing temperature and pressure resists contraction;

Important parameters: thermal conduction, cooling

macroscopic behavior depends on microscopic physics

Gravitational Instability



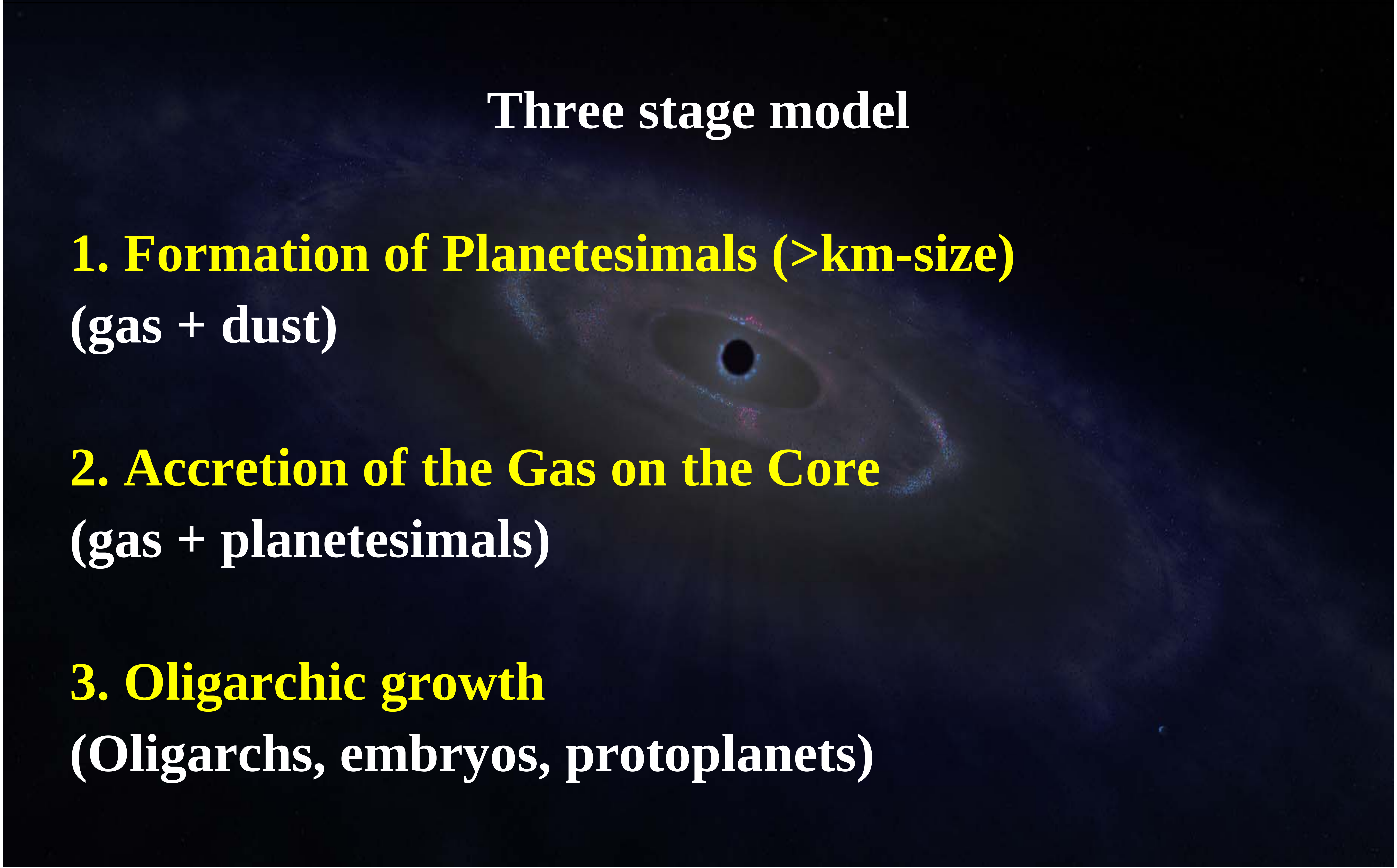
Problems:

Self-gravity: High mass protoplanetary disks
Result: Giant planets (earth?)
Radia: >50AU

Requriement: thermal conductivity - unrealistic

Turbulence

Core Accretion



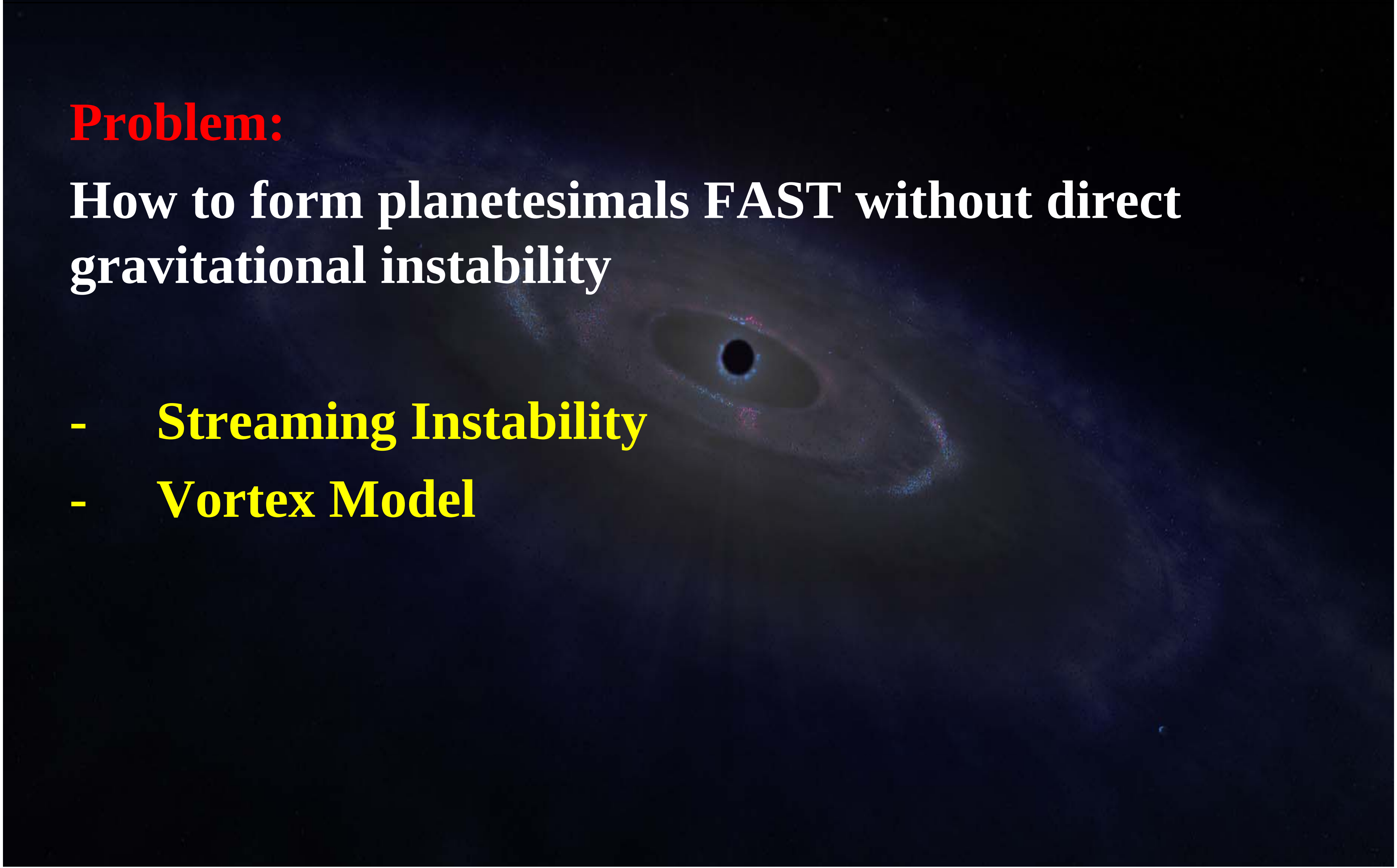
Three stage model

1. Formation of Planetesimals (>km-size)
(gas + dust)

2. Accretion of the Gas on the Core
(gas + planetesimals)

3. Oligarchic growth
(Oligarchs, embryos, protoplanets)

Core Accretion

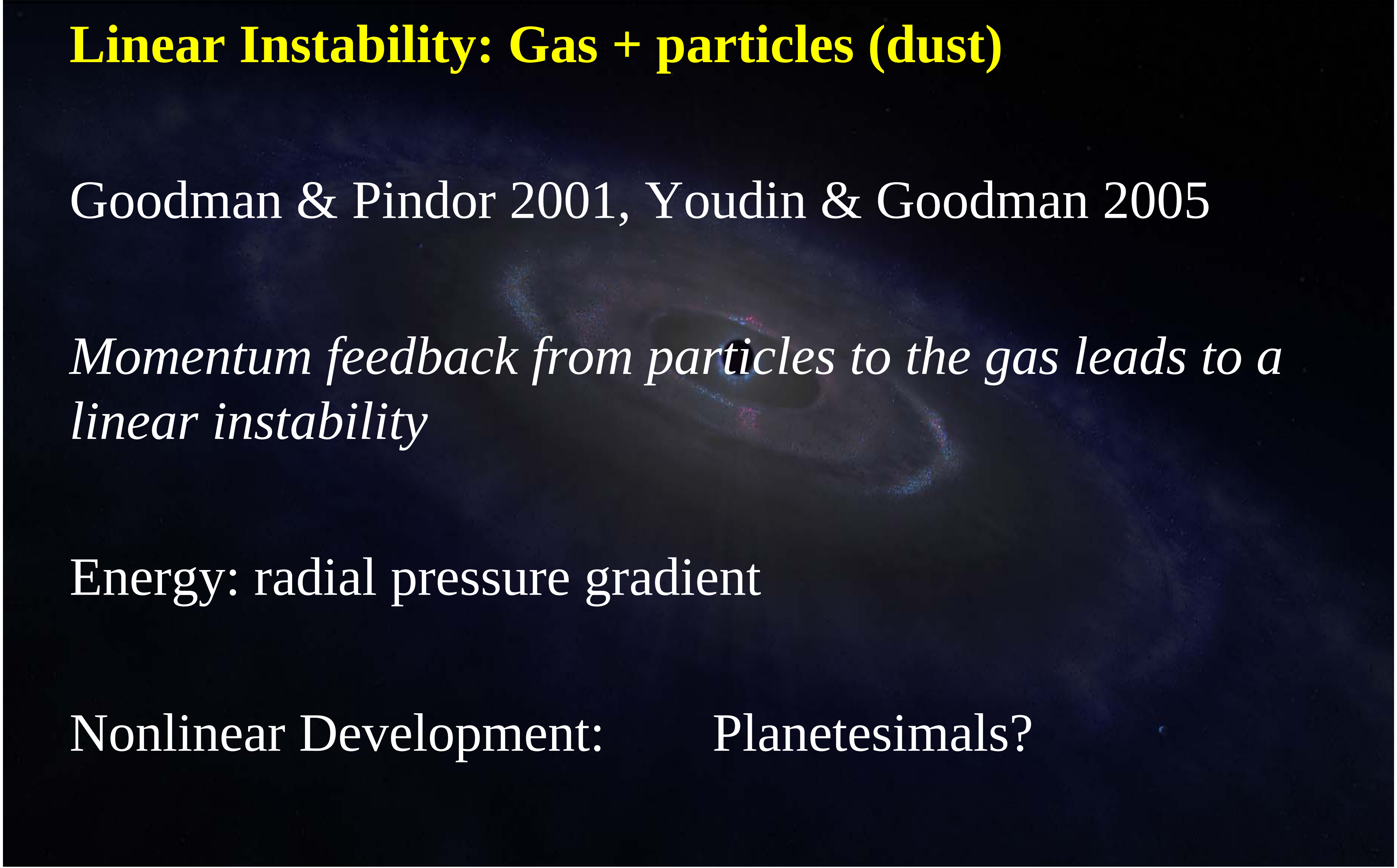


Problem:

How to form planetesimals FAST without direct gravitational instability

- **Streaming Instability**
- **Vortex Model**

Streaming Instability



Linear Instability: Gas + particles (dust)

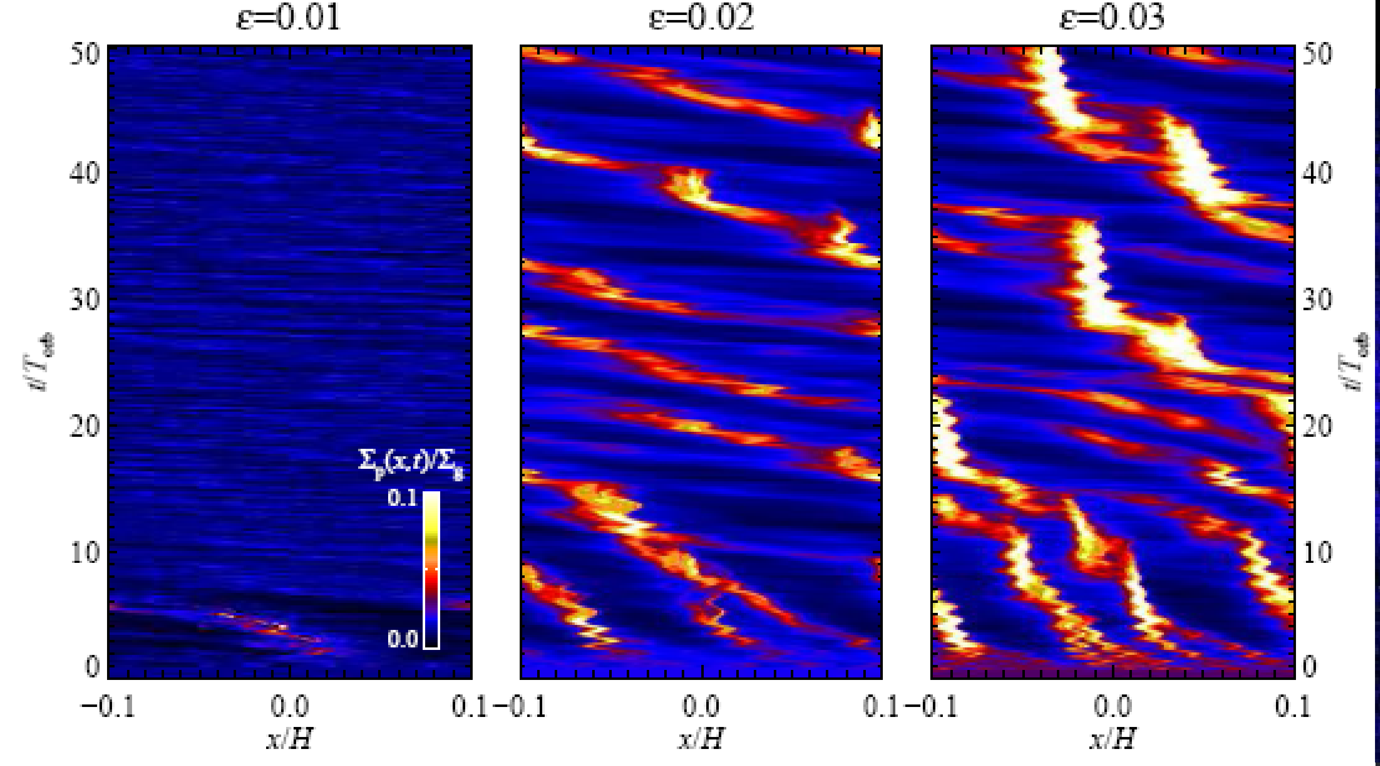
Goodman & Pindor 2001, Youdin & Goodman 2005

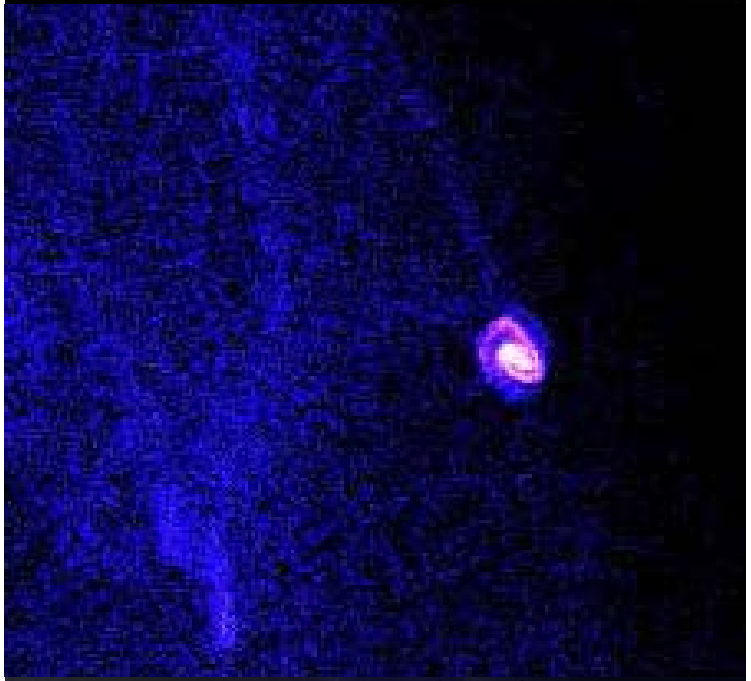
Momentum feedback from particles to the gas leads to a linear instability

Energy: radial pressure gradient

Nonlinear Development: Planetesimals?

Streaming Instability





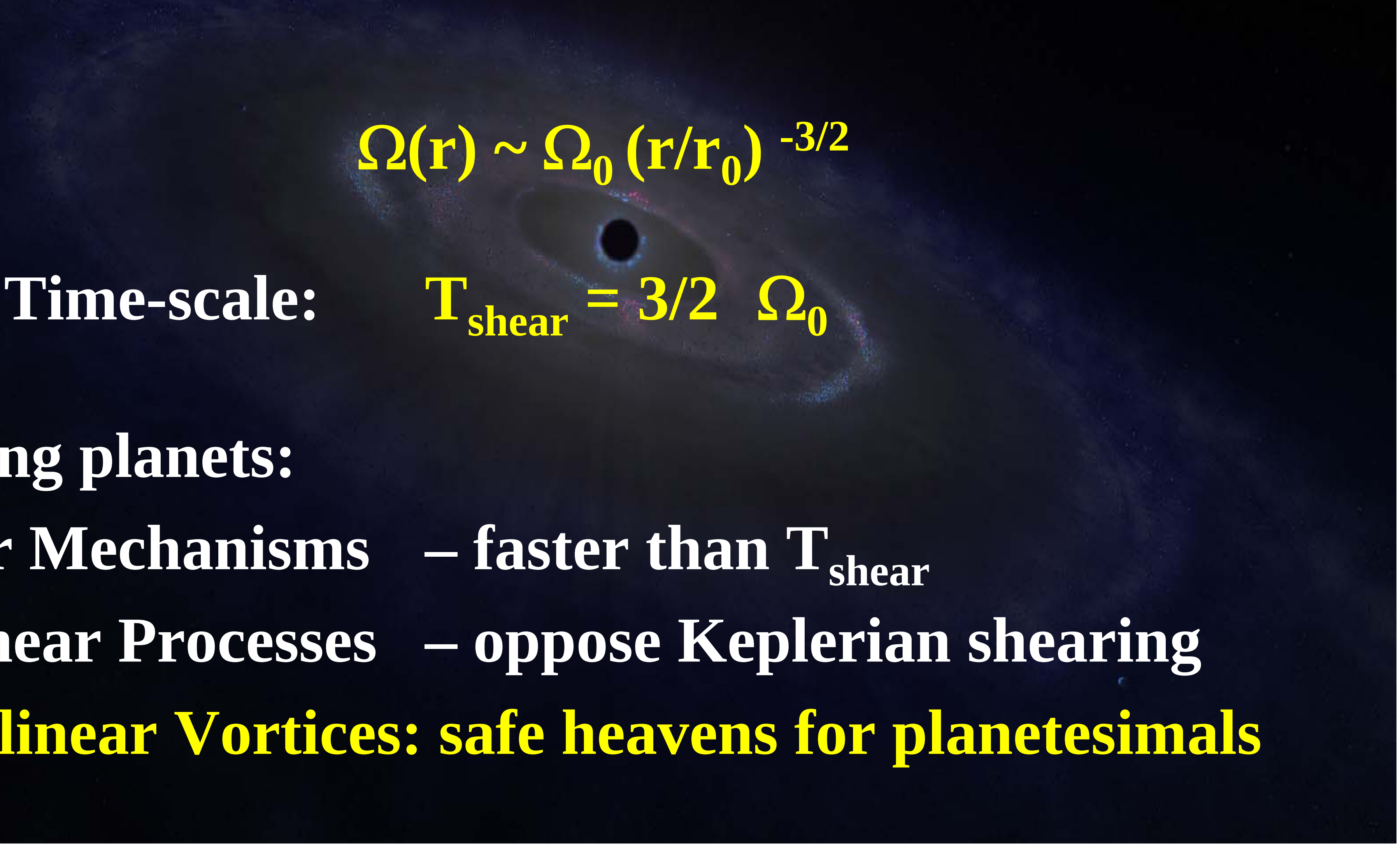
Turbulence: MRI? *Accelerates process (numerical)*

Problem: **Gas / Dust ratio**

Core Accretion: Vortex Model

Problem building structure in Keplerian Flows:

Strong local velocity shear (differential rotation)



$$\Omega(r) \sim \Omega_0\,(r/r_0)^{-3/2}$$

Shear Time-scale: $T_{shear} = 3/2\ \ \Omega_0$

Building planets:

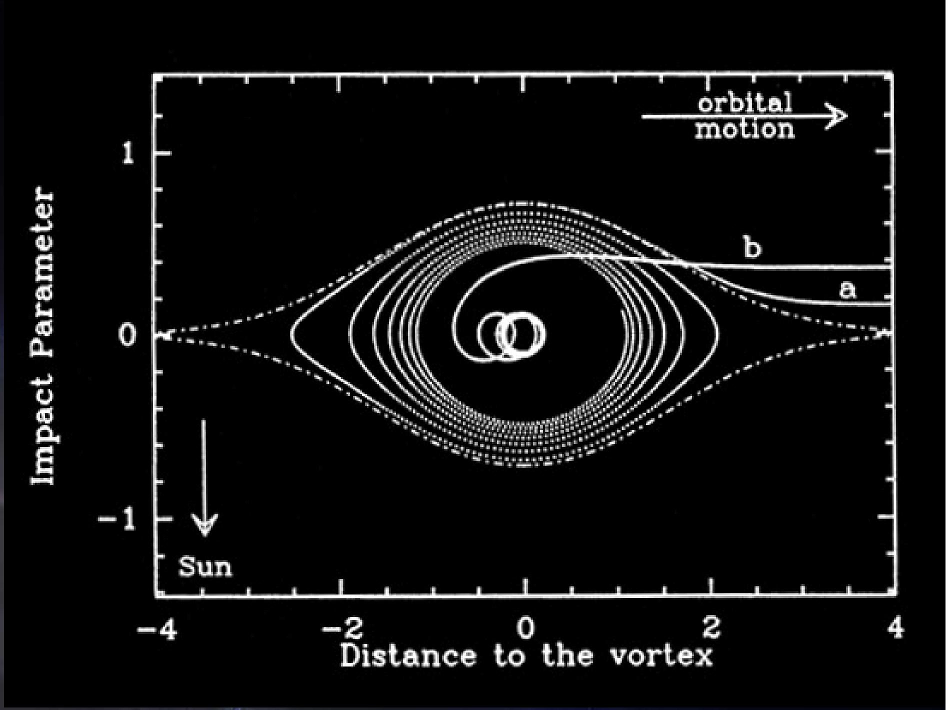
Linear Mechanisms – faster than T_{shear}

Nonlinear Processes – oppose Keplerian shearing

Nonlinear Vortices: safe heavens for planetesimals

Core Accretion: Vortex Model

Barge Sommeria (1995)



Dust capture by
long-lived vortices

- confirmed by numerical simulations

Long lived anticyclonic vortices can kinematically TRAP dust

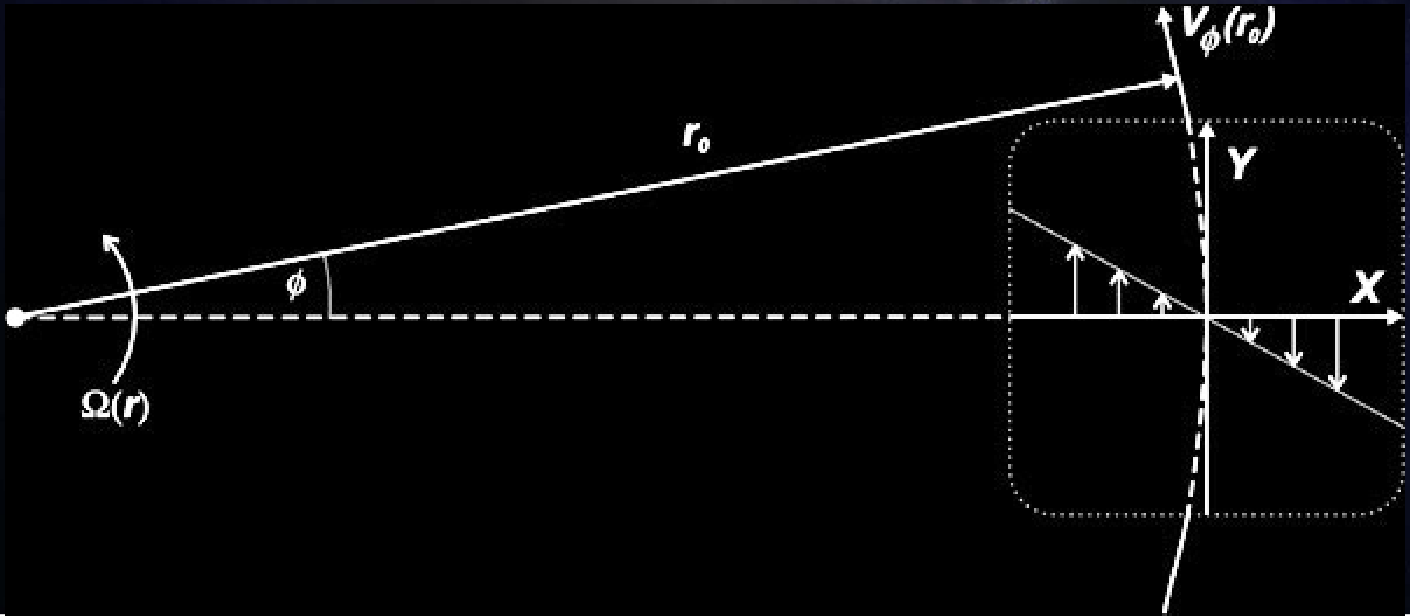
Anticyclonic Vortices

Stable vortex configuration in Keplerian flows

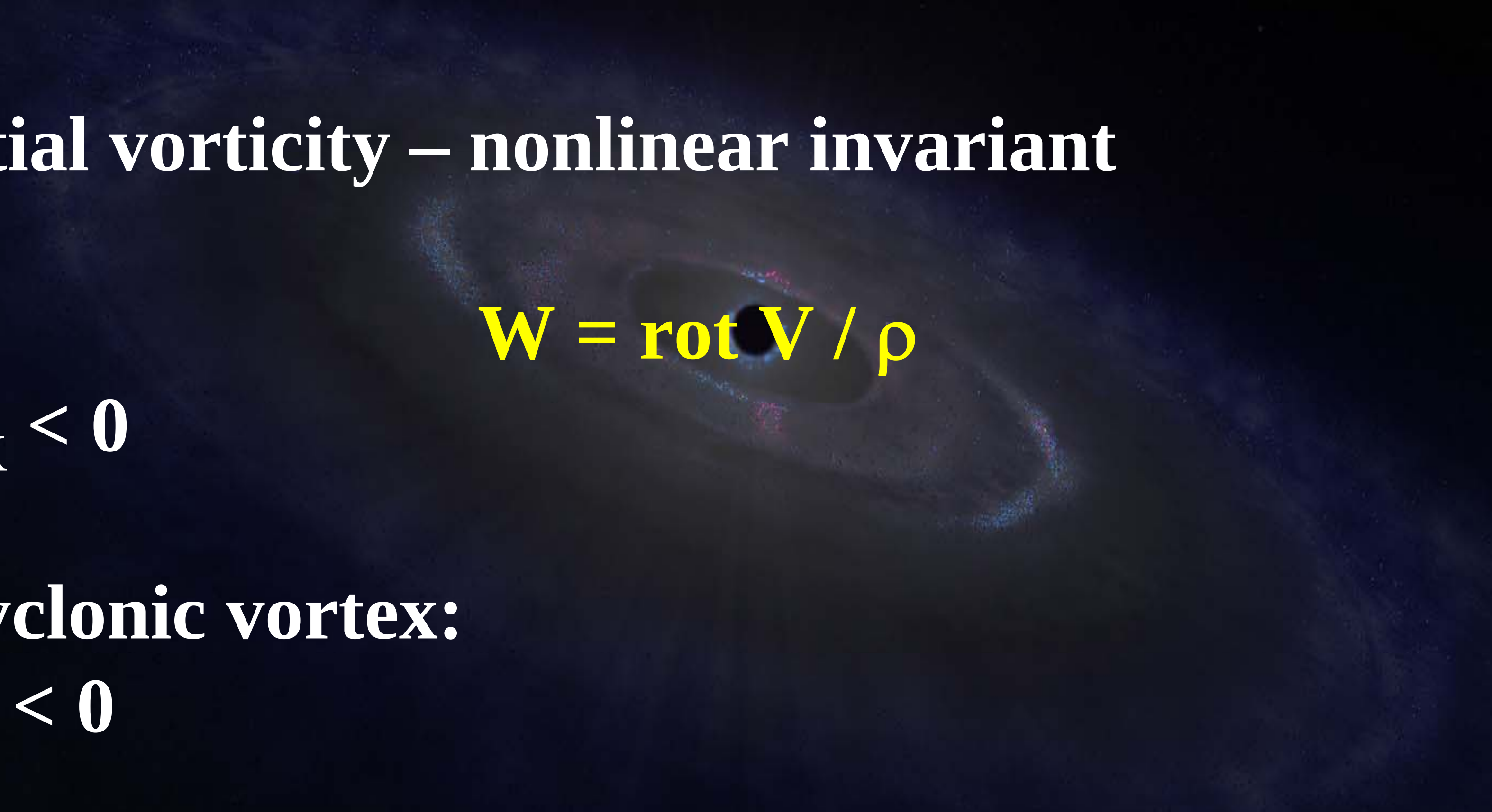
$V(r) \sim r^{-1/2}$

$d\,\Omega(r) / d\,r = -3/2\,\Omega(r)$

direction of vortex rotation is opposite to global circulation



Anticyclonic Vortices



Keplerian flow: vortical

Potential vorticity – nonlinear invariant

$$W = \mathrm{rot}\, V / \rho$$

$\mathrm{rot}\, V_K < 0$

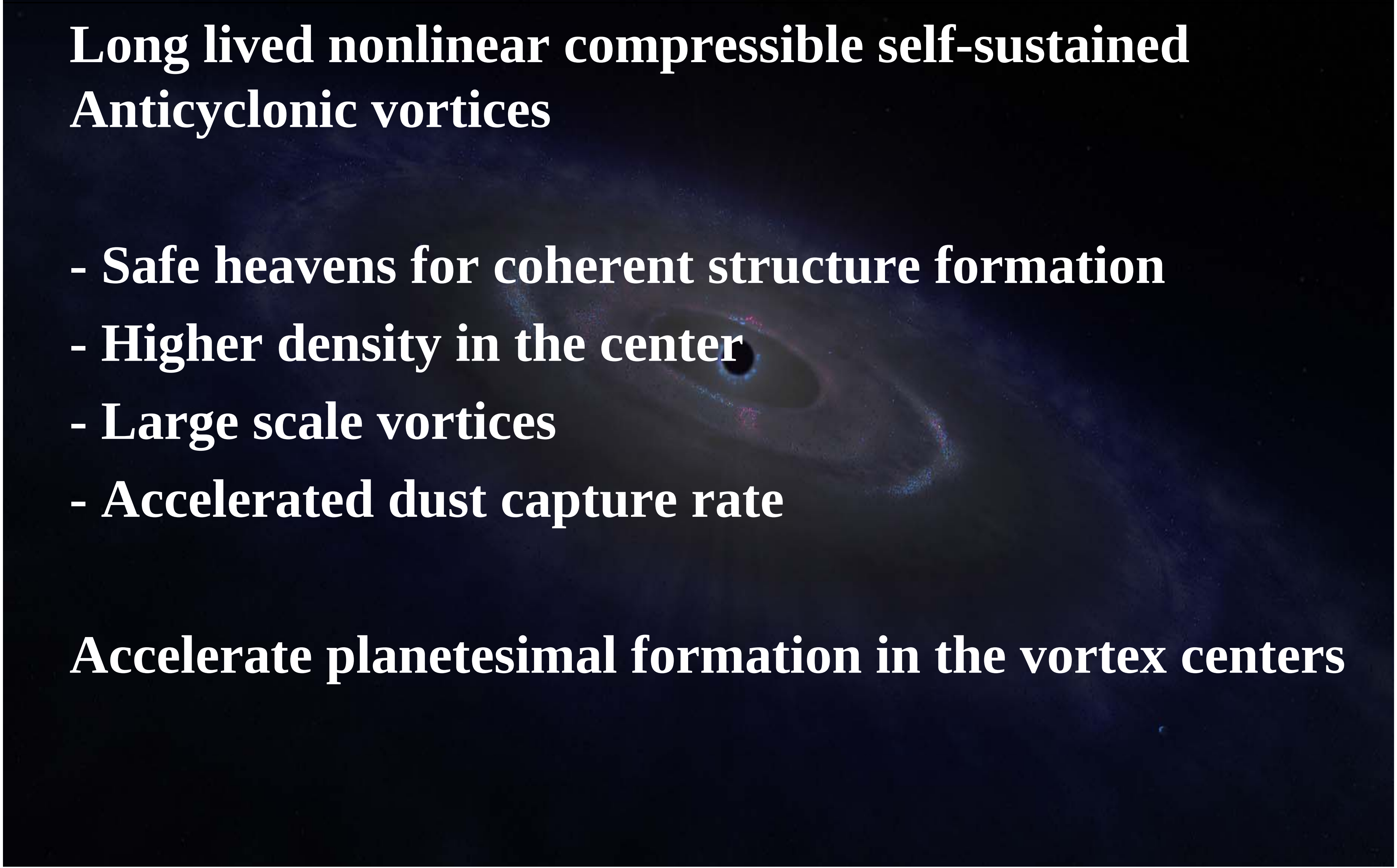
Anticyclonic vortex:

$\mathrm{rot}\, V_1 < 0$

$\rho 1 > 0$

Maximal density area in the vortex center

Vortices in Keplerian Flows



Long lived nonlinear compressible self-sustained Anticyclonic vortices

- Safe heavens for coherent structure formation

- Higher density in the center

- Large scale vortices

- Accelerated dust capture rate

Accelerate planetesimal formation in the vortex centers

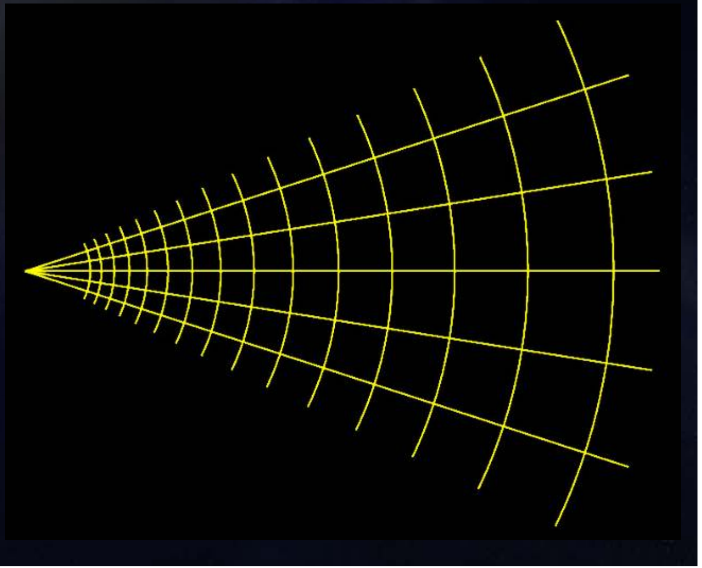
Vortex Stability Simulations

DNS of vortex dynamics in HD Keplerian disks

(Bodo et al. 2007)

code: **PLUTO** (plutocode.to.astro.it)
solver: **Riemann/Godunov, HD, FARGO, (ppm)**
grid: **Polar, [8192x1559]**

- Global compressible model
- Radially inhomogeneous grid
- Shock capturing method



Vortex Stability Simulations

Initially imposed:

Cyclonic and anticyclonic vortices with different amplitudes and size

Non-equilibrium distribution of potential vorticity

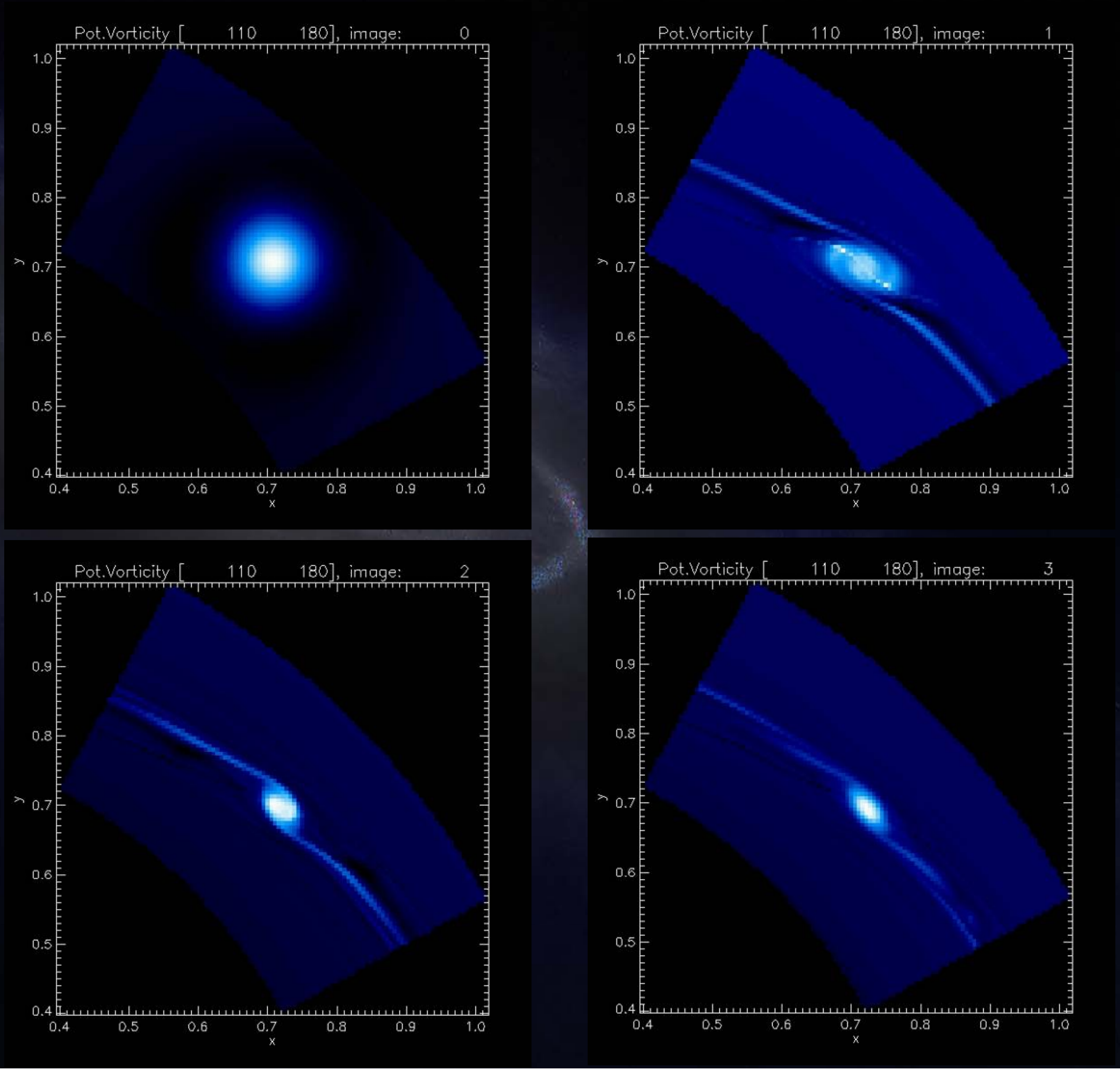
Nonlinear developments:

- **Direct nonlinear adjustment to stable vortex;**
- **Stable vortex configuration**
- **Long time evolution (compressibility, dissipation)**

Vortex Adjustment

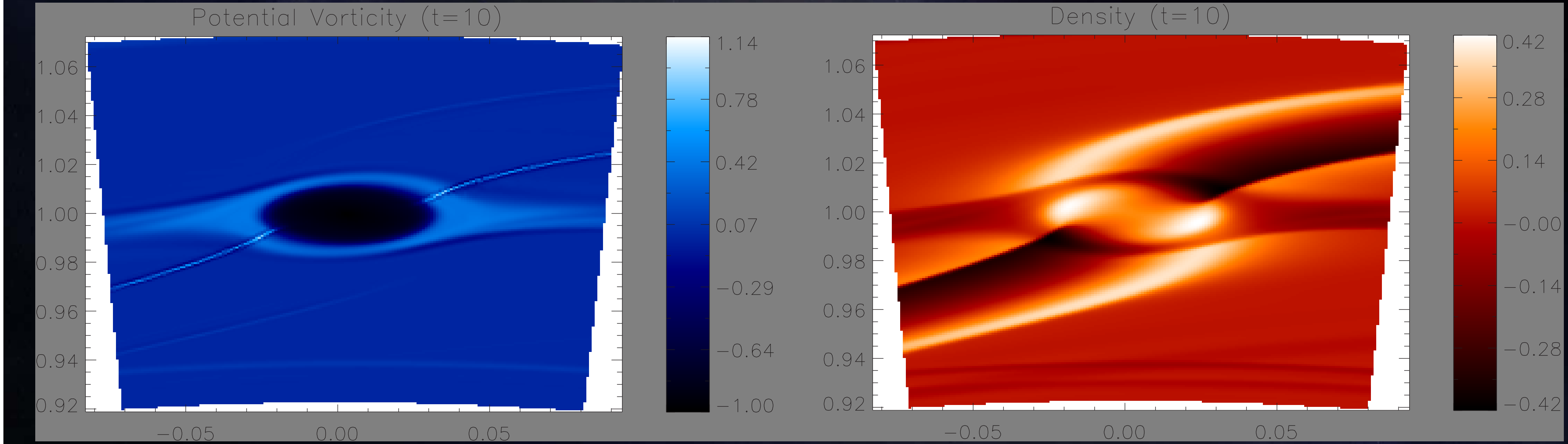
Adjustment of potential vorticity

Time scale: ~3-5 Ω^{-1}



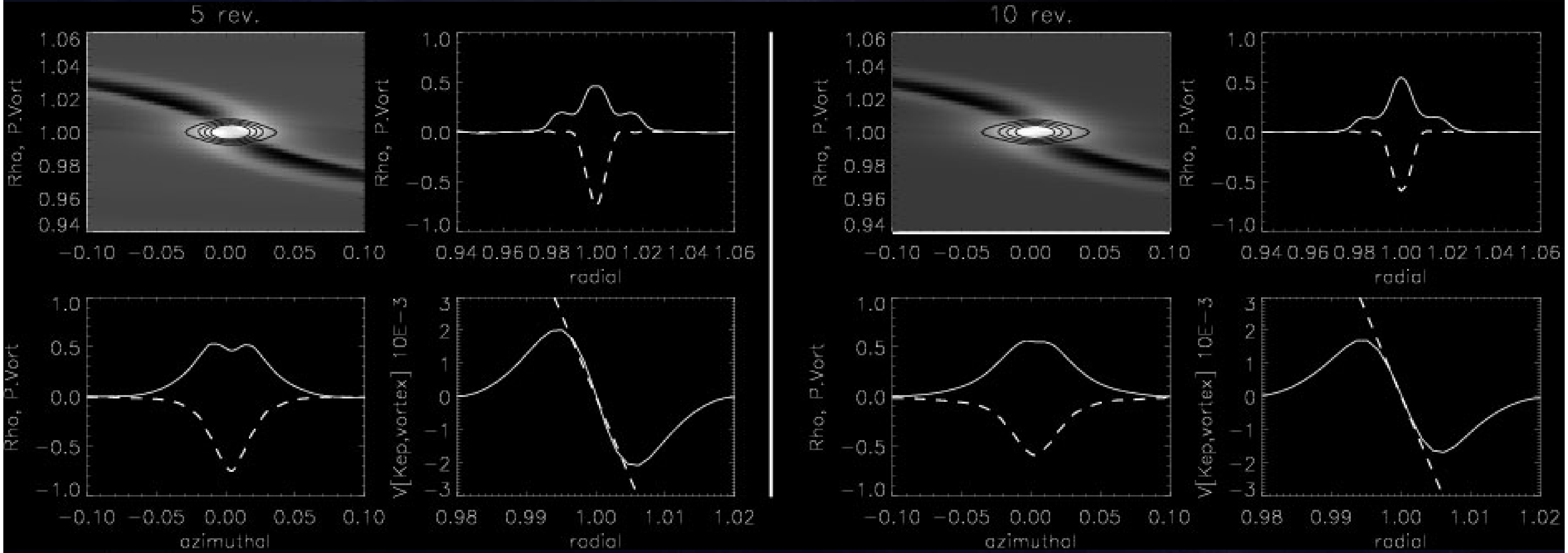
Vortex Structure

Configuration of the long-lived anticyclonic vortex



Vortex Structure

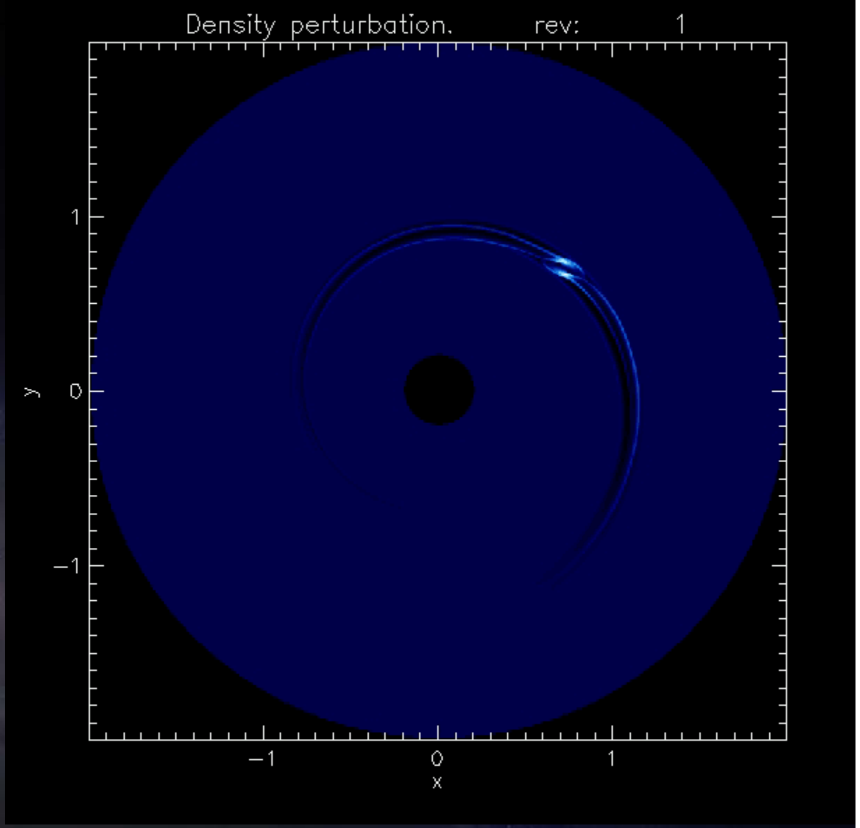
Time evolution

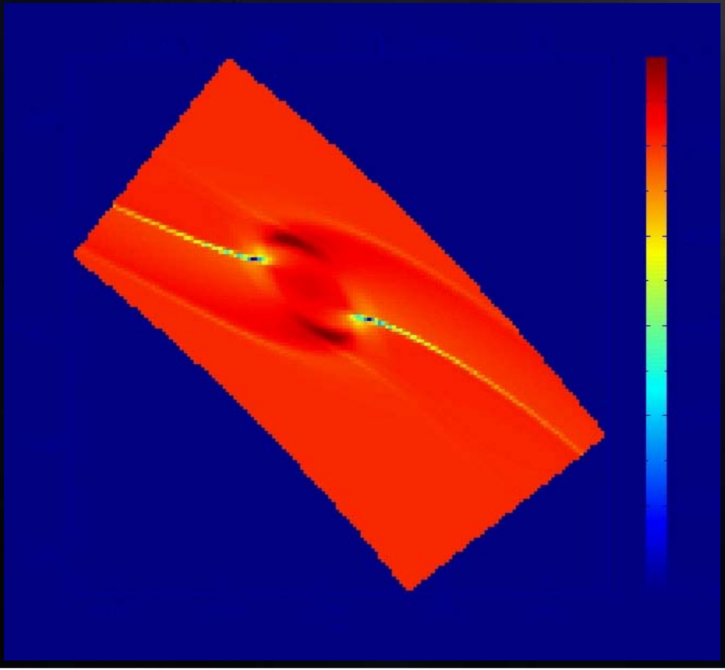


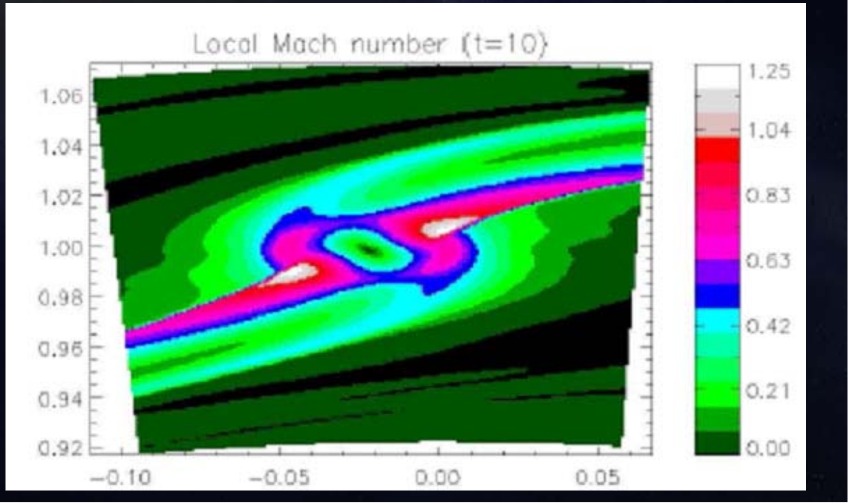
Spiral Shocks

Steady pattern of shock waves

Shocks induced by vortex without embedded planetesimal

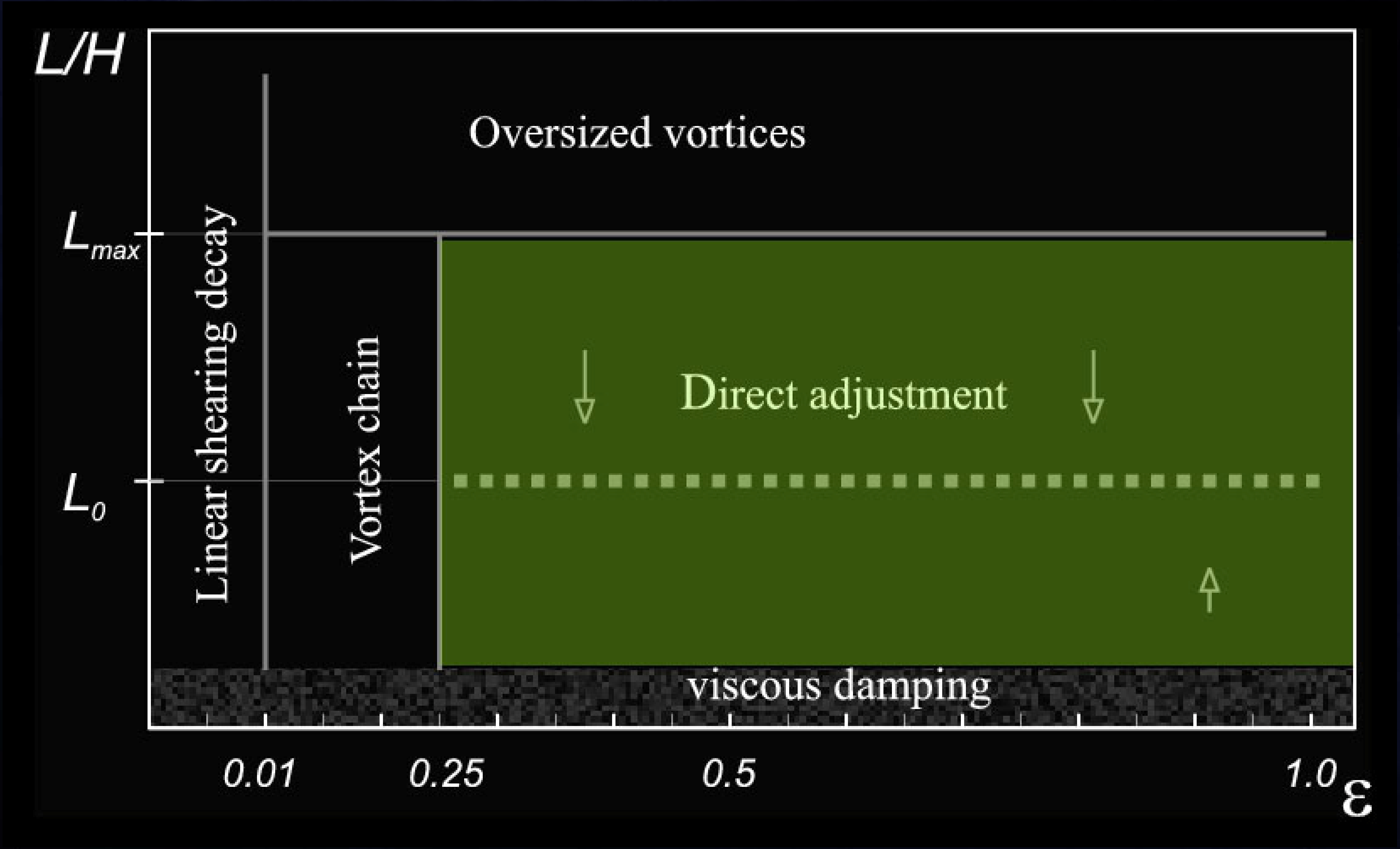






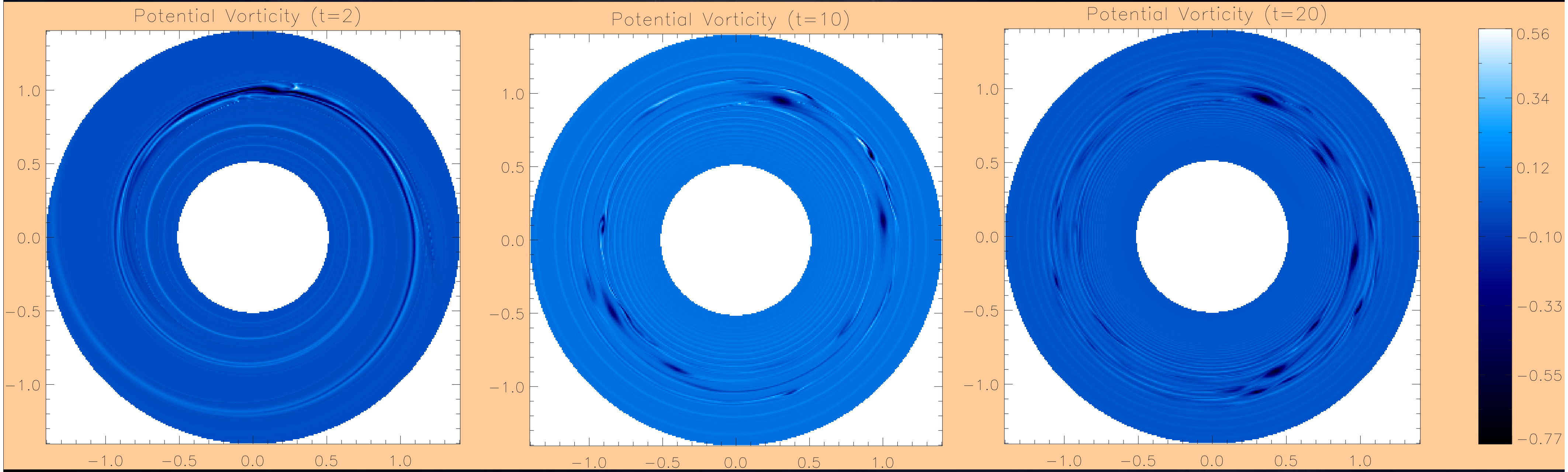
Nonlinear Adjustment

- Nonlinear thresholds in PV amplitude
- Size constraints



Oversized Vortices

Radial Spreading due to Rossby waves

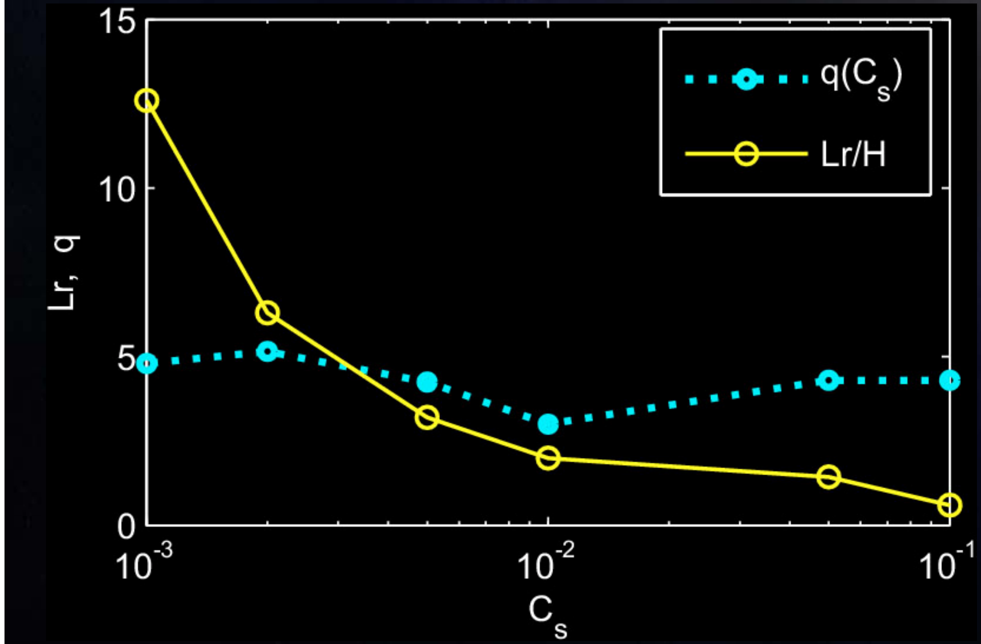


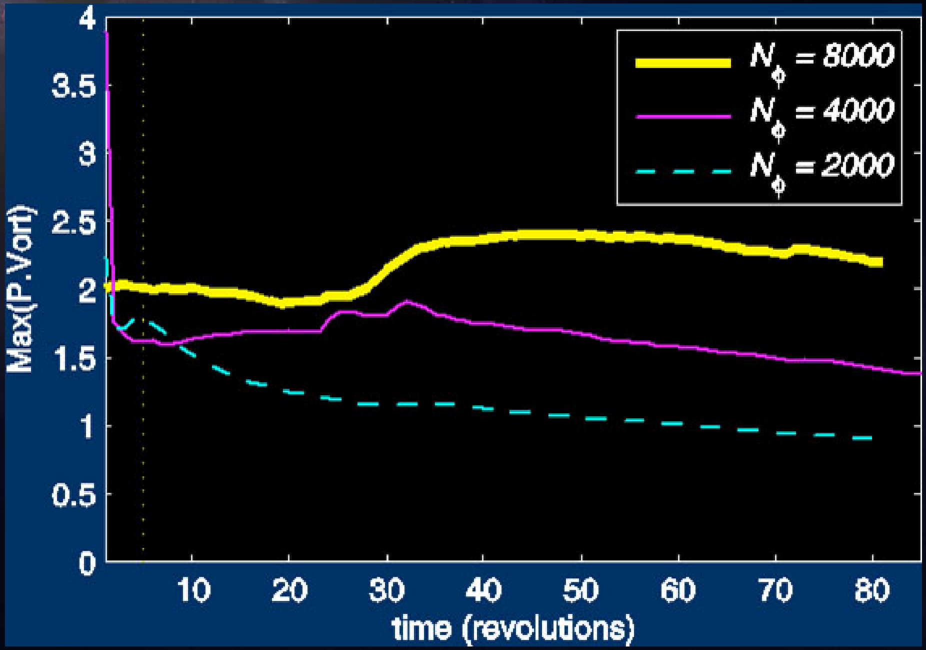
Equilibrium Vortices

Scaling relations of self-sustained vortex:

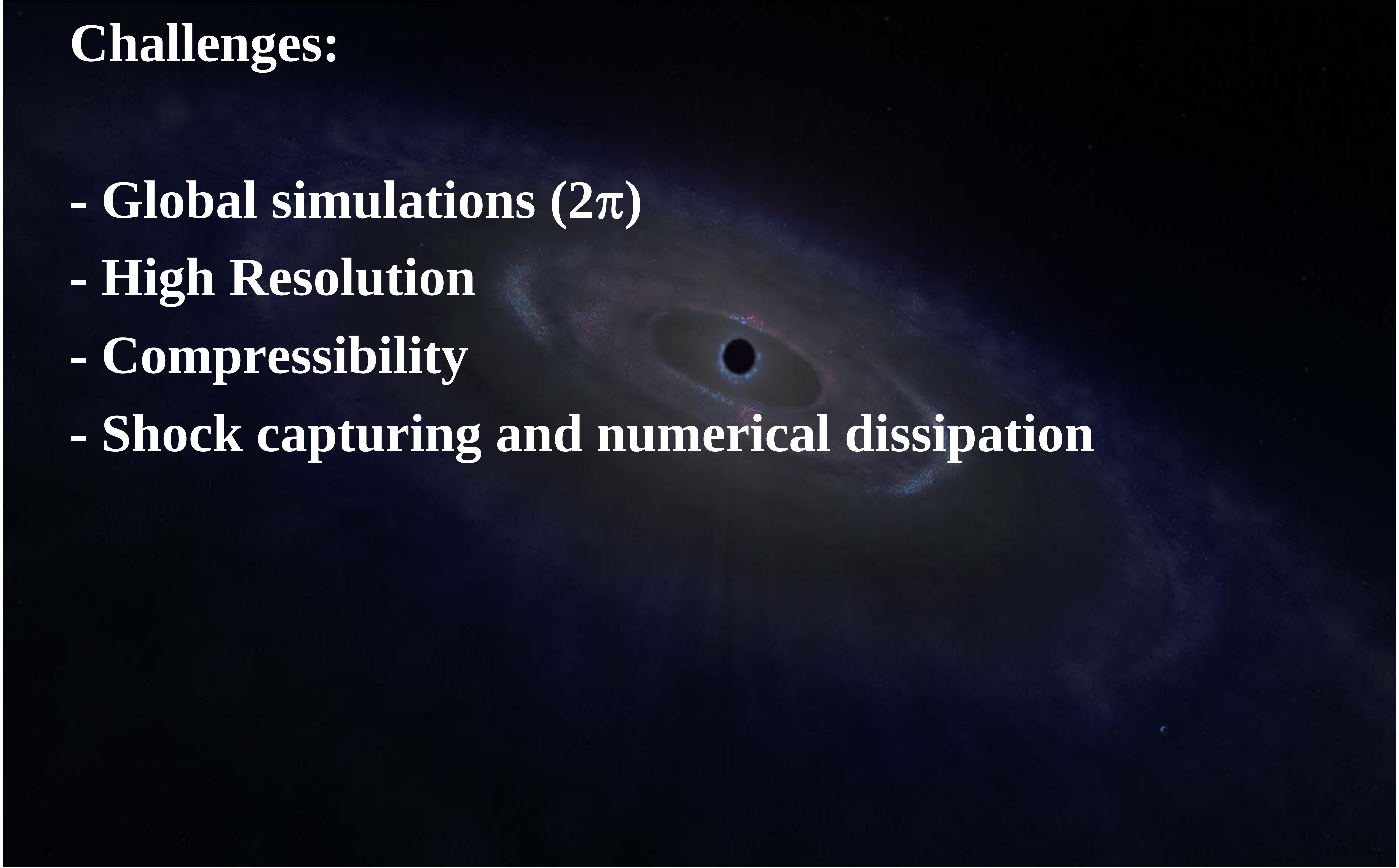
a = f(Cs)

q=5





Vortex Simulations



Challenges:

- Global simulations (2π)
- High Resolution
- Compressibility
- Shock capturing and numerical dissipation

3D Elliptic Instability

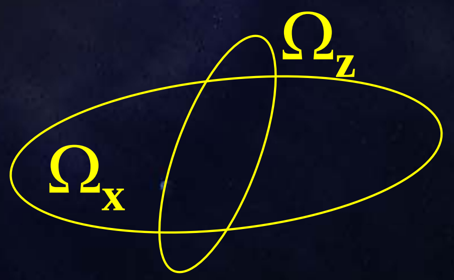
Bayly (1986), Pierrehumbert (1986)

Instability of 3D vortices

Internal parametric resonance between vertical and horizontal circular motions

RESULT:

destruction of coherent structure



Elliptic Instability

Anticyclonic vs Cyclonic vortices

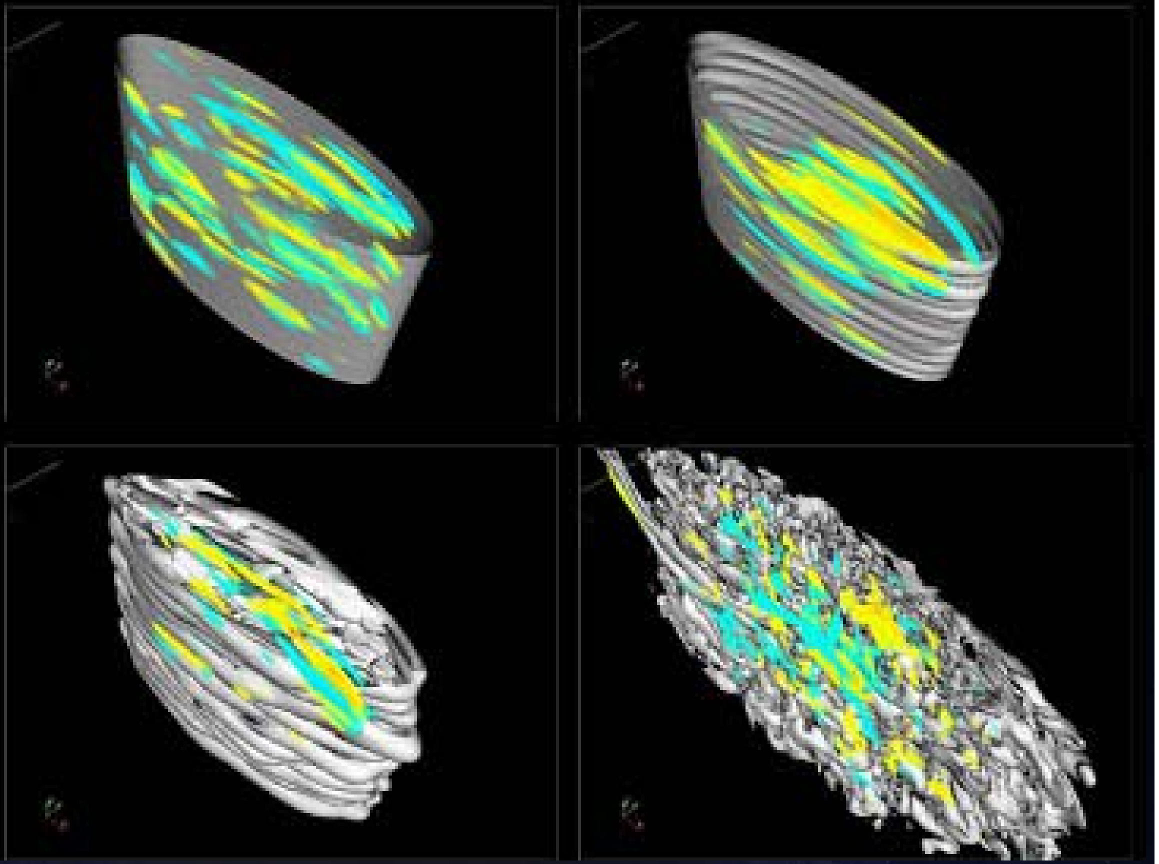
Laboratory
Experiments

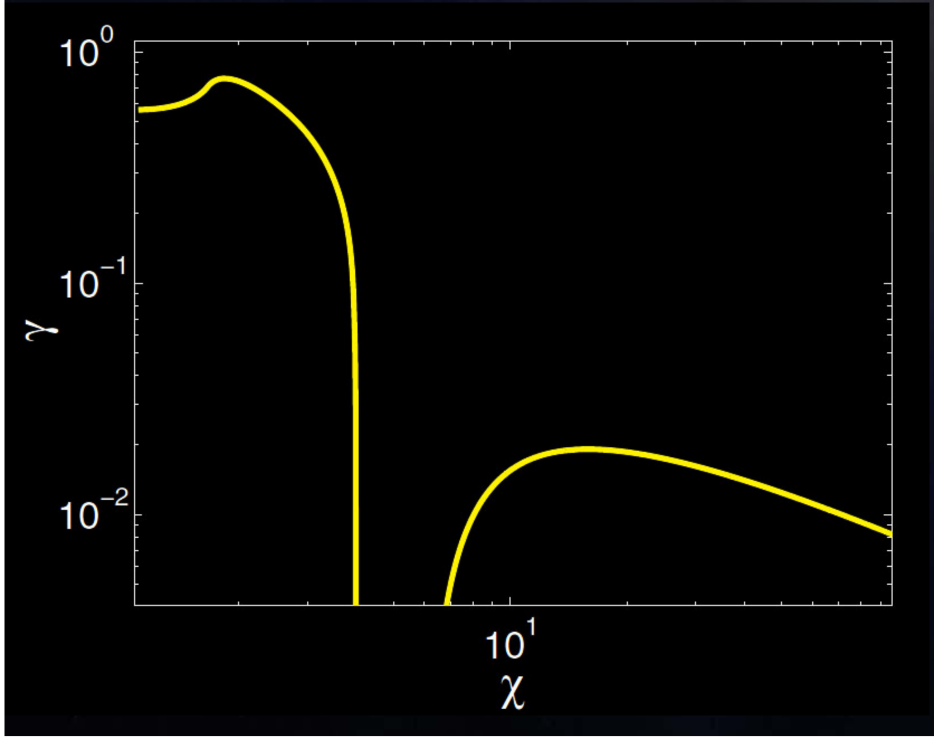
Instability:
anticyclonic circulation

Afanasyev (2006)

Vortex break down

Lesur Papaloizou (2009)

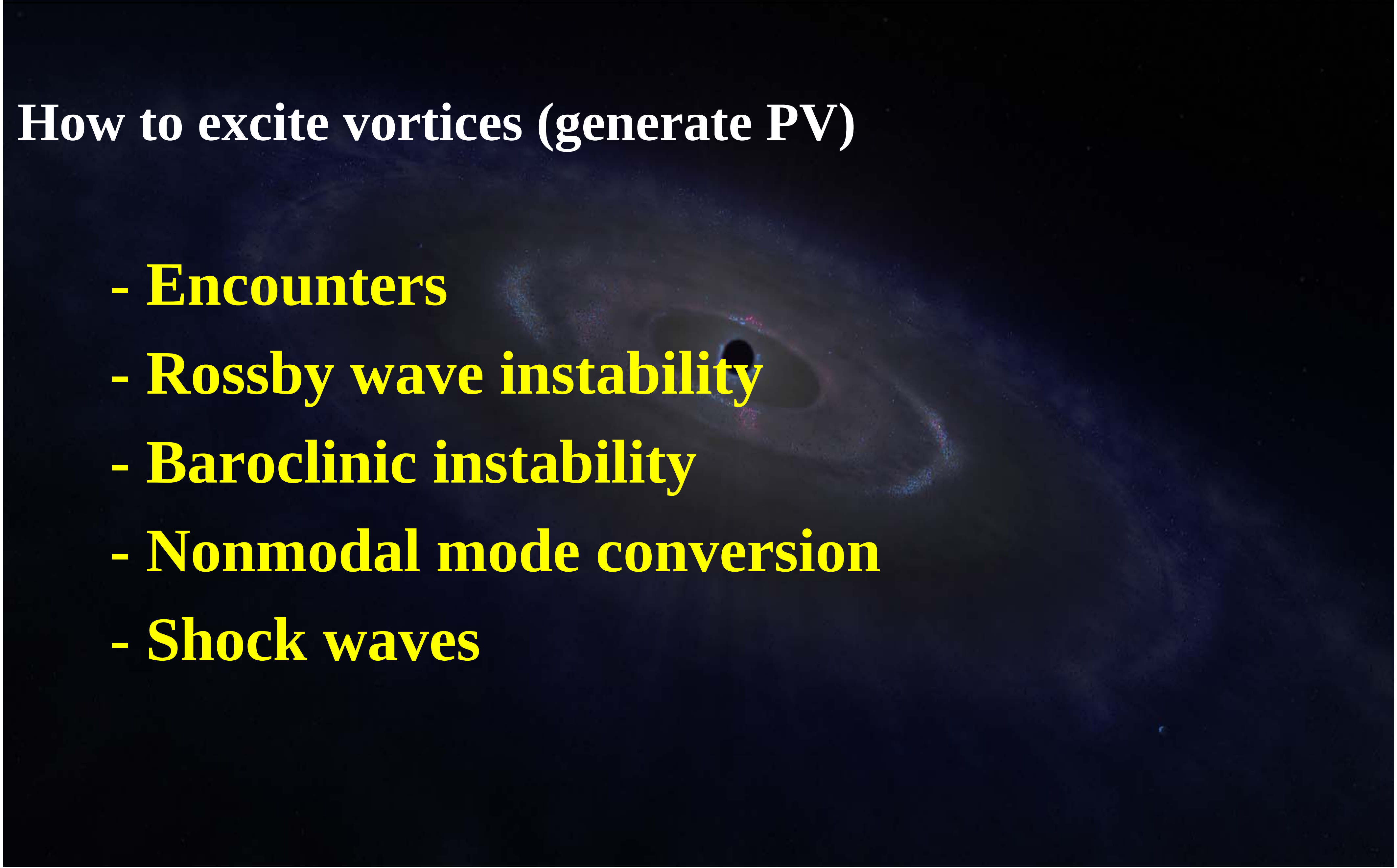




No instability at:

$4 < q < 5.9$

Vortex Sources



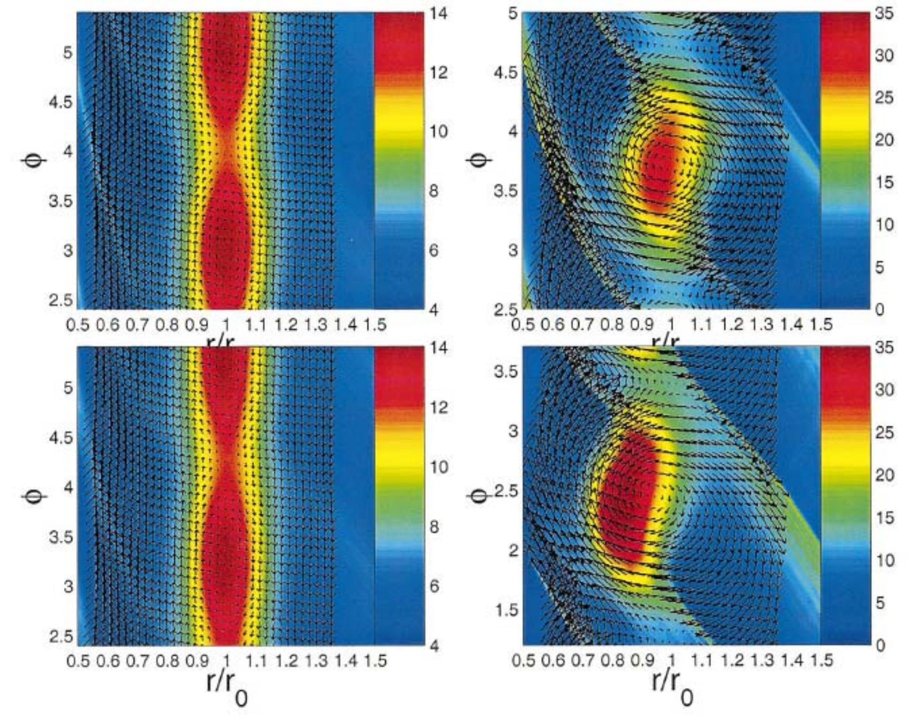
How to excite vortices (generate PV)

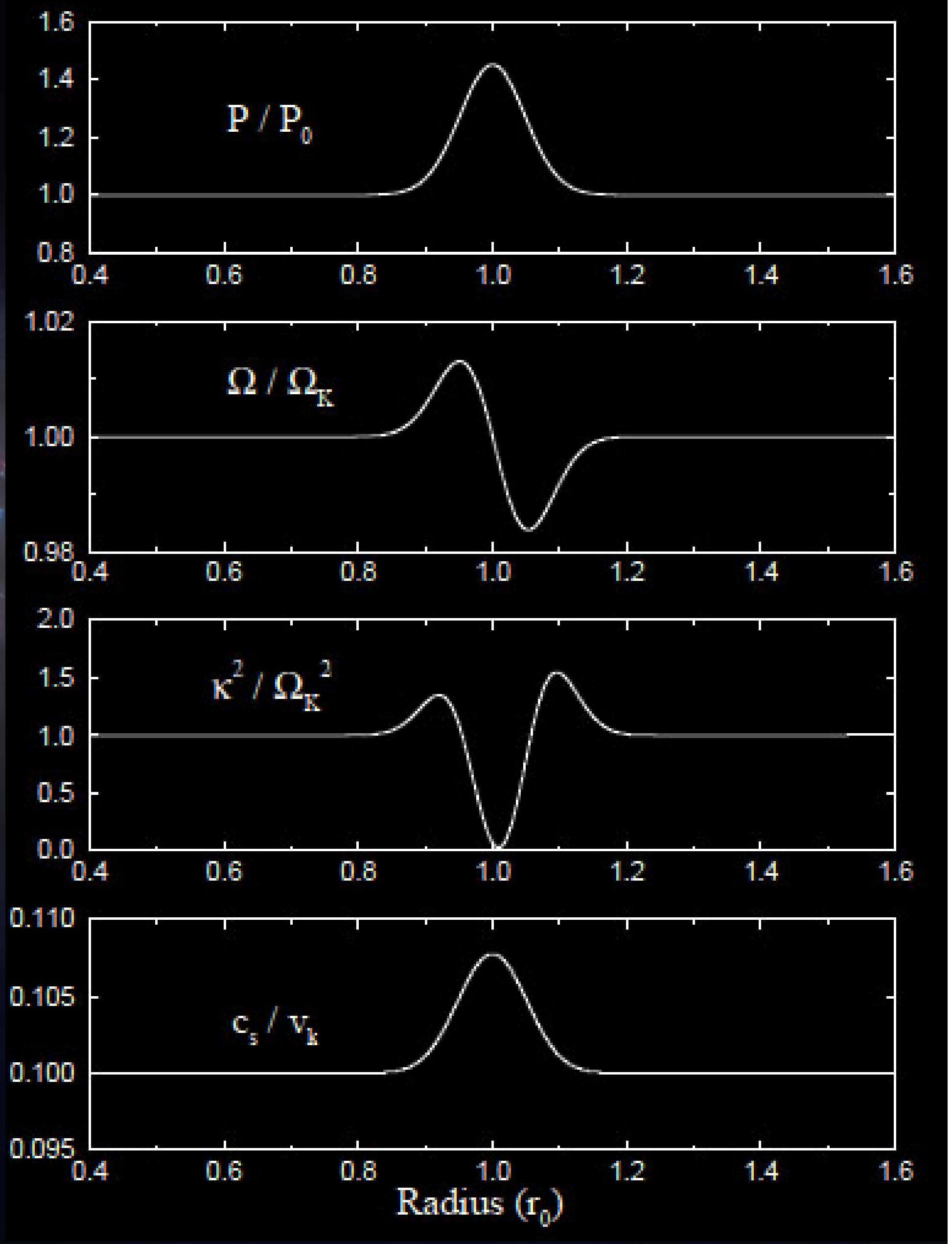
- Encounters
- Rossby wave instability
- Baroclinic instability
- Nonmodal mode conversion
- Shock waves

Rossby Wave Instability

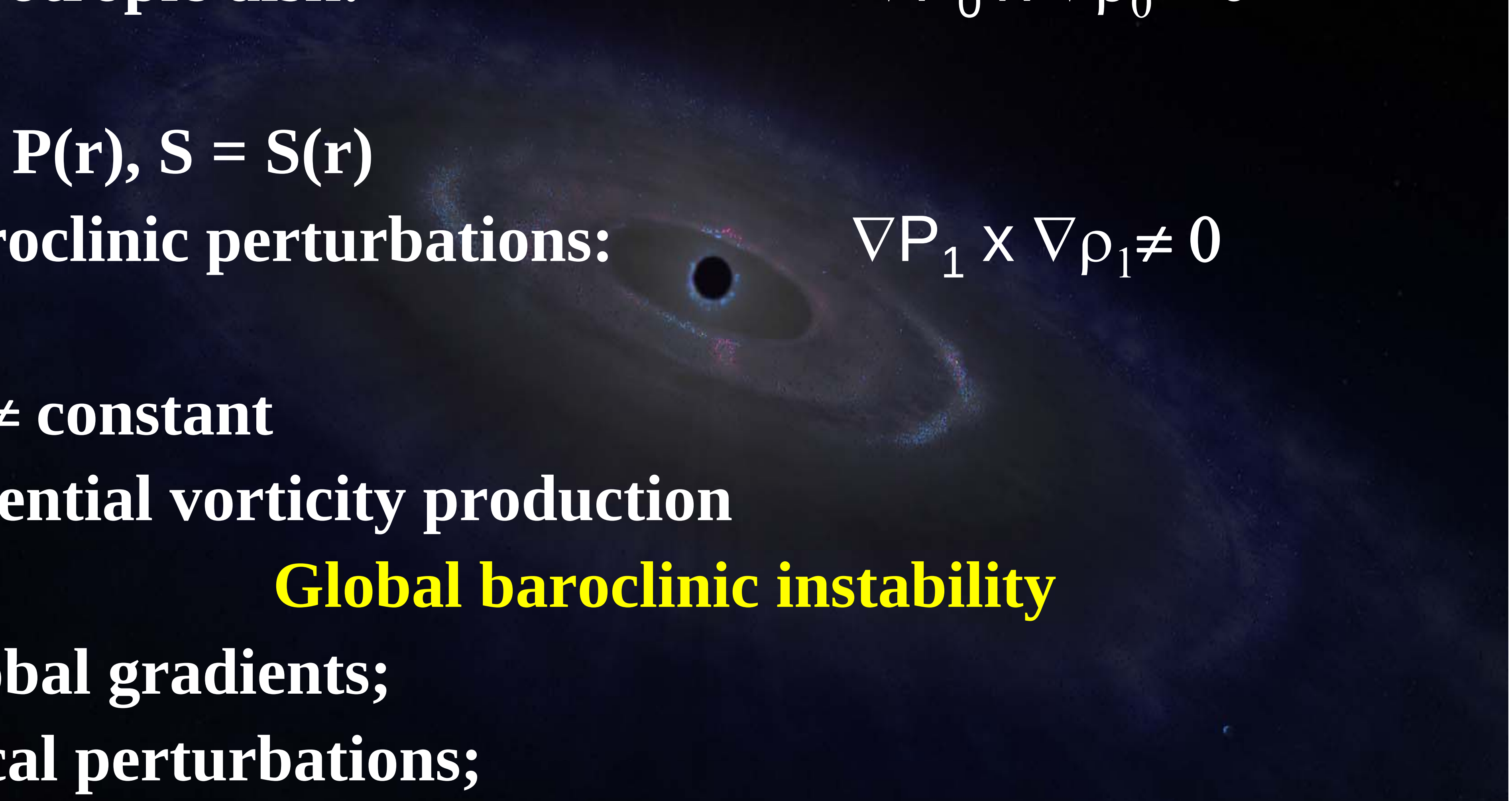
(Colgate, Lovelace 1999)

Thermodynamic bump
in the disk structure





Baroclinic Instability



Barotropic disk: $\nabla P_0 \times \nabla \rho_0 = 0$

P = P(r), S = S(r)

Baroclinic perturbations: $\nabla P_1 \times \nabla \rho_1 \neq 0$

W ≠ constant

Potential vorticity production

Global baroclinic instability

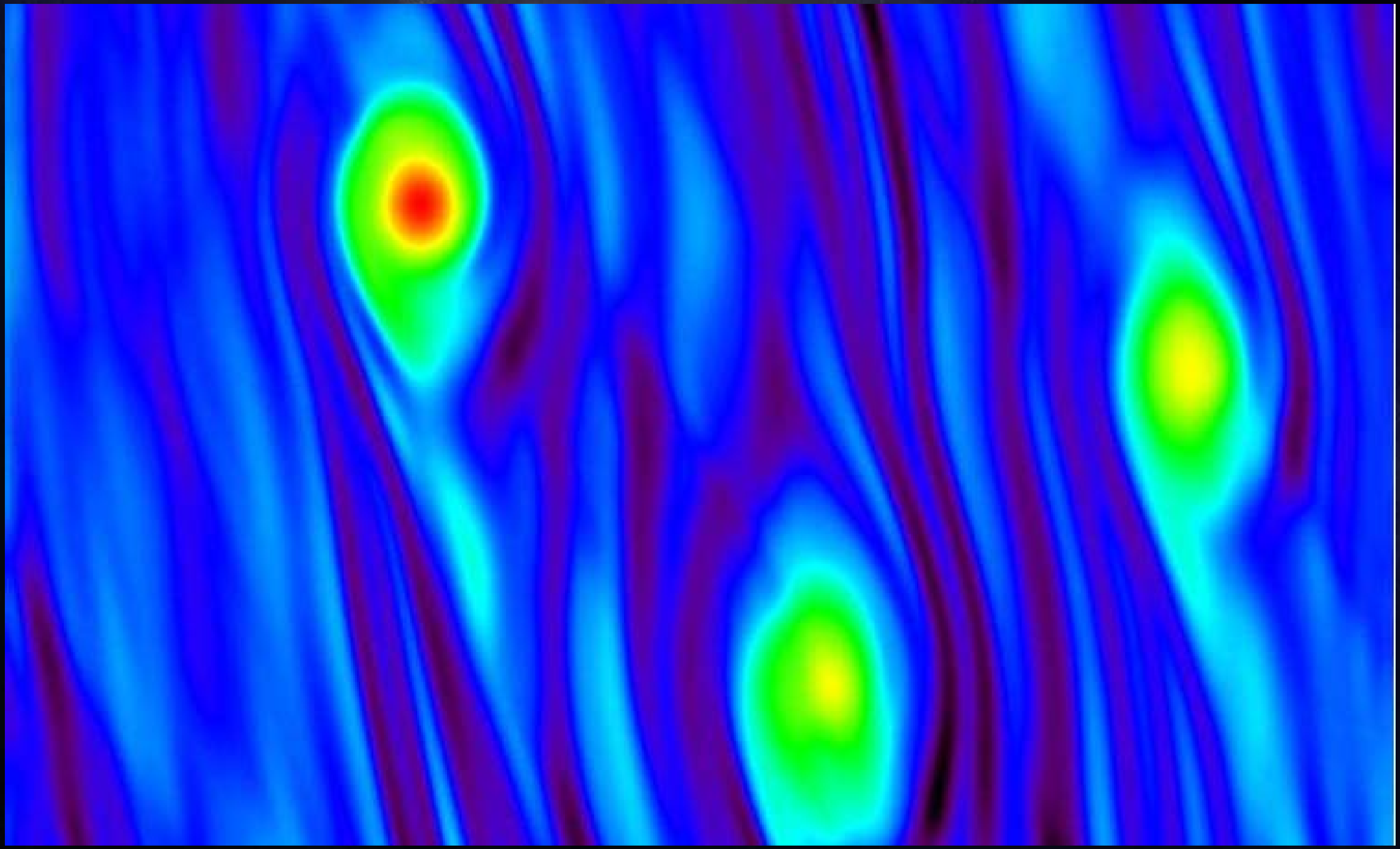
Global gradients;

Local perturbations;

Baroclinic Instability

(Klahr, Bodenheimer 2003)

Thermal perturbations lead to development of anticyclonic vortices in disks with global stratification



Baroclinic Instability

Radial stratification in protoplanetary disk

$$\bar{\Sigma}(r) = \Sigma_0 \left(\frac{r}{r_0}\right)^{-\beta_\Sigma} , \quad \bar{P}(r) = P_0 \left(\frac{r}{r_0}\right)^{-\beta_P}$$

Time-scale of the growth of baroclinic instability

$$T_{baro} \sim \gamma^2 / (\beta_P \, \beta_S)$$

Observations: $\beta_S < 1$

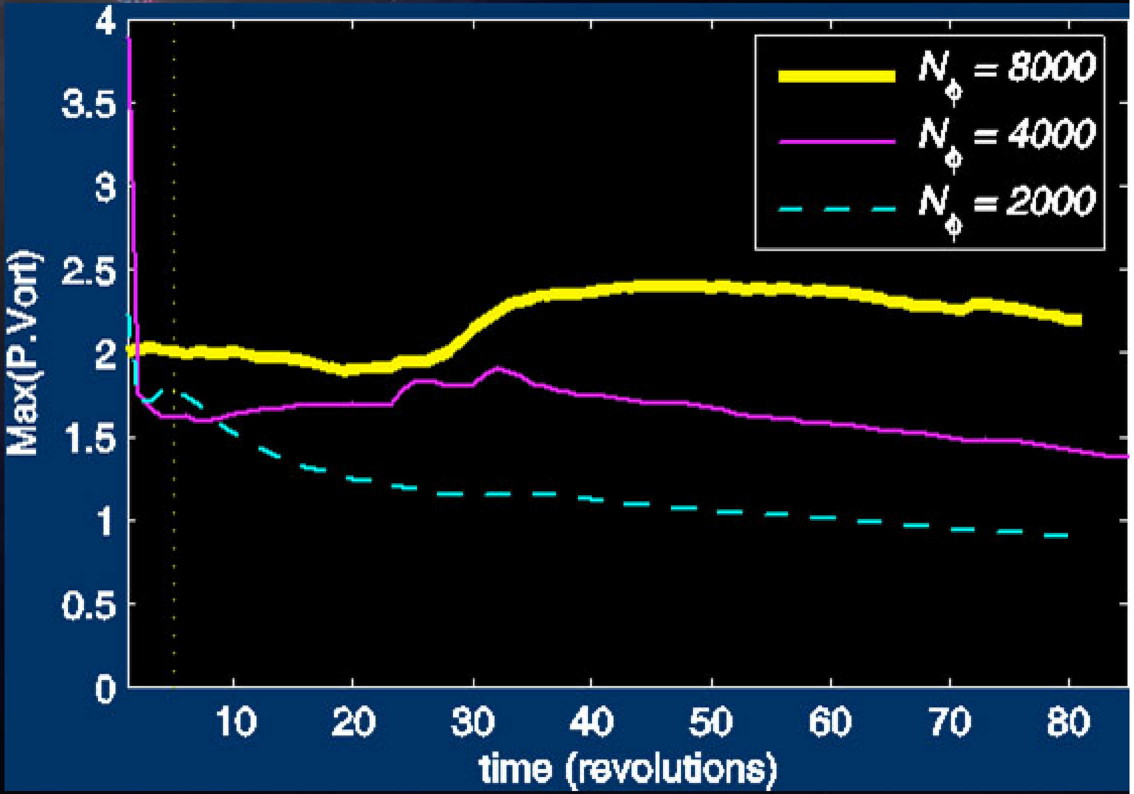
$$T\Omega > 100$$

Baroclinic growth is much slower than Keplerian shearing

PV production

PV evolution at high resolution runs:
damping and growth

Nonlinear production of
potential vorticity:
spiral shock waves

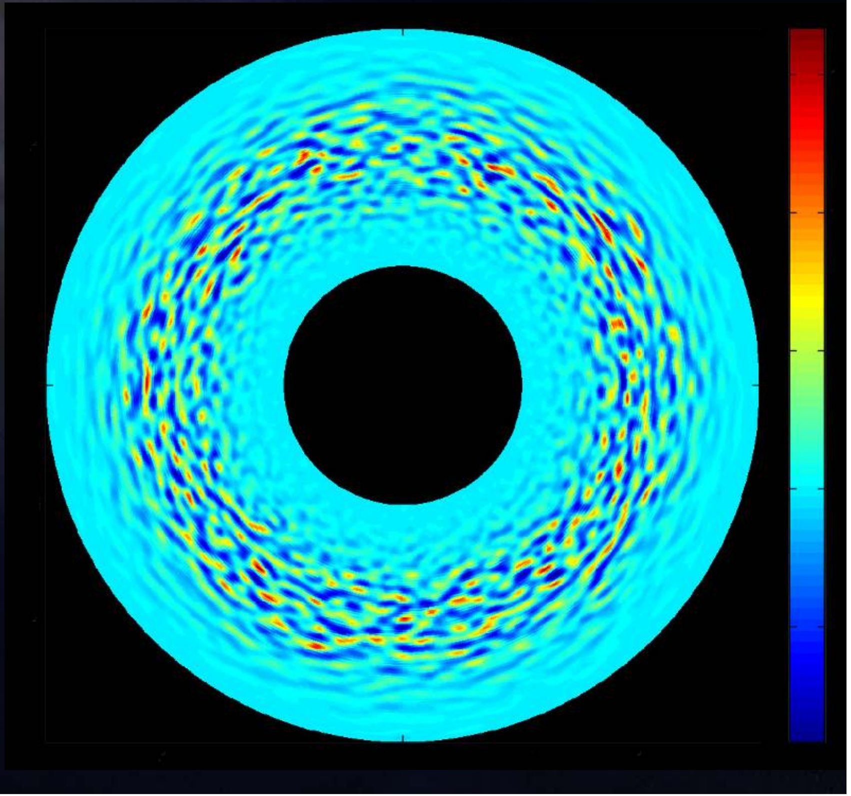


Simulating Pressure Perturbations

Stochastic pressure perturbations

$$p(r,\varphi) = \exp\left\{ -\left(\frac{(r-R_0)^2}{\Delta R_0{}^2} \right)^n \right\} \Theta(r,\varphi)$$

- No entropy
- No potential vorticity



Radial profile

Spectrum

Gaussian noise

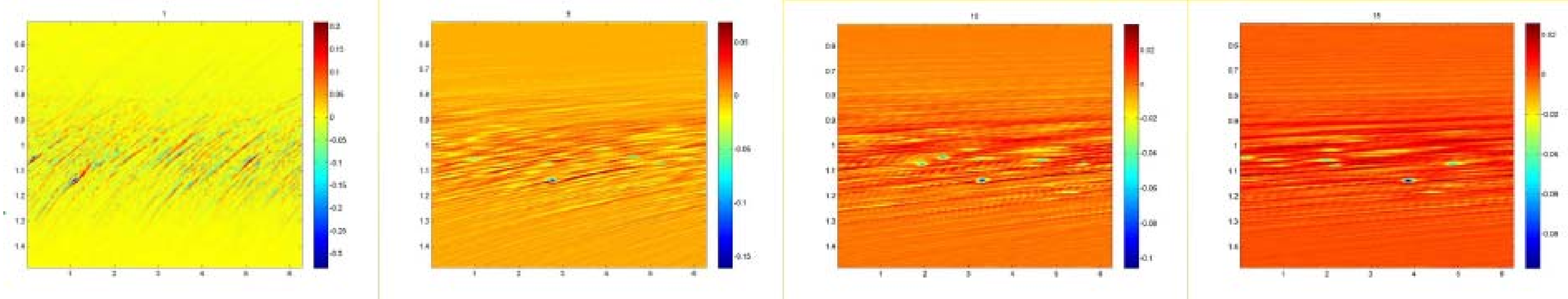
PV Shock Production

Shock waves (~ 1 local disk revolution)

Vorticity (~2 local disk revolutions)

Strong Anticyclonic vortices

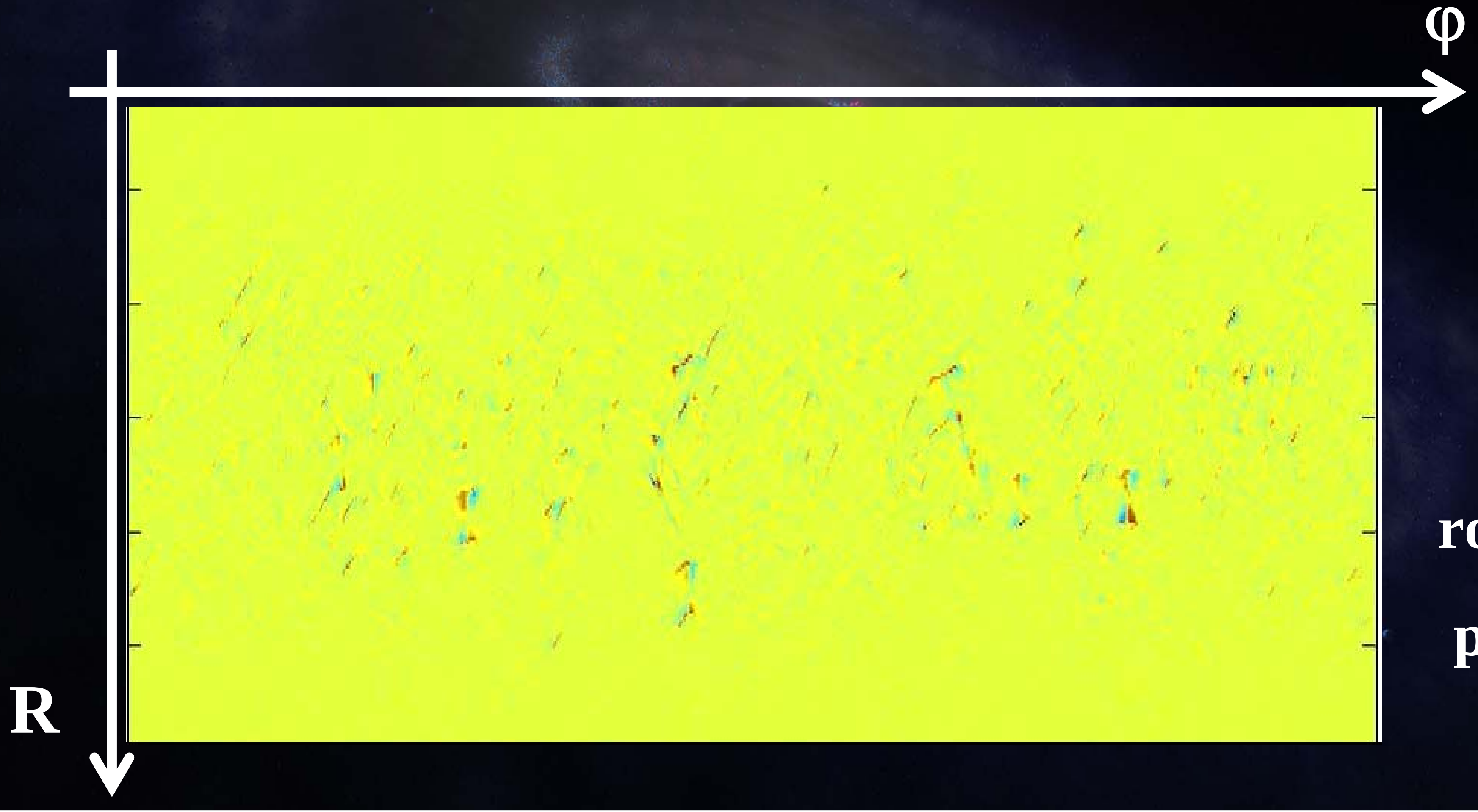
(~5 local disk revolutions)



Numerical Simulations

Excitation of potential vorticity and development of stable anti-cyclonic vortices

Global frame (2D Cartesian visualization)

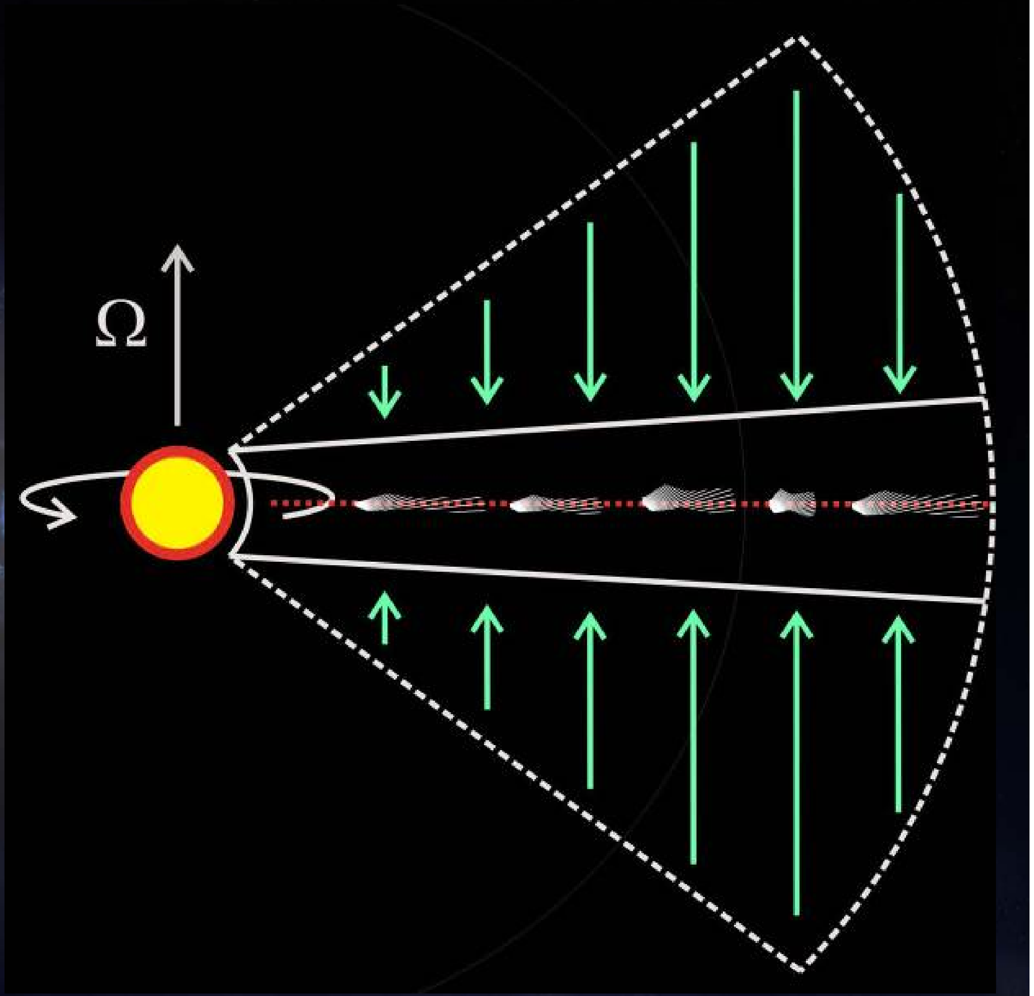


20

rotation

periods

Formation of Protoplanetary Disks



Disk flattering

middle layer:

+ compressible perturbations

+ Shock waves

Initial Heating

(Dullemond 2009)

PV shock production

PV production on shearing time scales

Keplerian flow:

2-3 strong self-sustained anticyclonic vortices

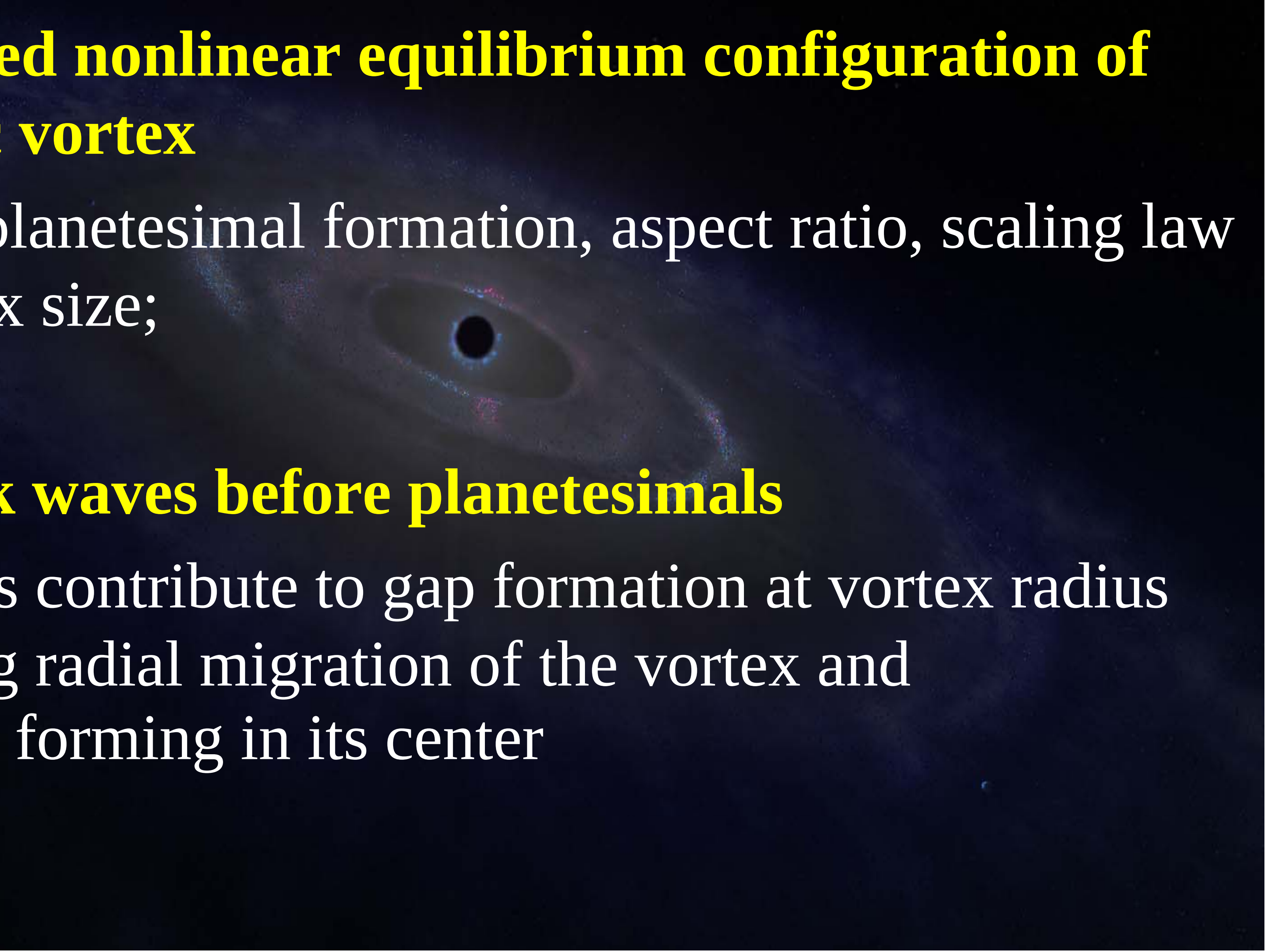
Sub-Keplerian flows:

Radial stratification increases number of vortices that survive nonlinear developments

Remains to be clarified:

effects of vertical boundary conditions

Summary



Self-sustained nonlinear equilibrium configuration of anticyclonic vortex

Centers for planetesimal formation, aspect ratio, scaling law for the vortex size;

Spiral shock waves before planetesimals

Spiral shocks contribute to gap formation at vortex radius thus stopping radial migration of the vortex and planetesimal forming in its center

Summary

PV excitation during initial flattering of the protoplanetary disk

Pressure perturbations develop into shocks that produce potential vorticity

Turbulent evolution of stochastic pv leads to development of several strong anticyclonic vortices

cyclonic vortices shear away and dissipate, while some anticyclonic vortices merge to produce stable structures

Summary

Pressure and density perturbations at early stages can play important role in vortex development

compressibility factor can be equally important for gravitational fragmentation, as well as initial stages of the core accretion model

Observations: primordial vorticity?

მადლობთ ყურადღებისთვის